Exploring the Dark Universe

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Physikalisches Kolloquium, Marburg, 08.02.2018

www.app.uni-freiburg.de

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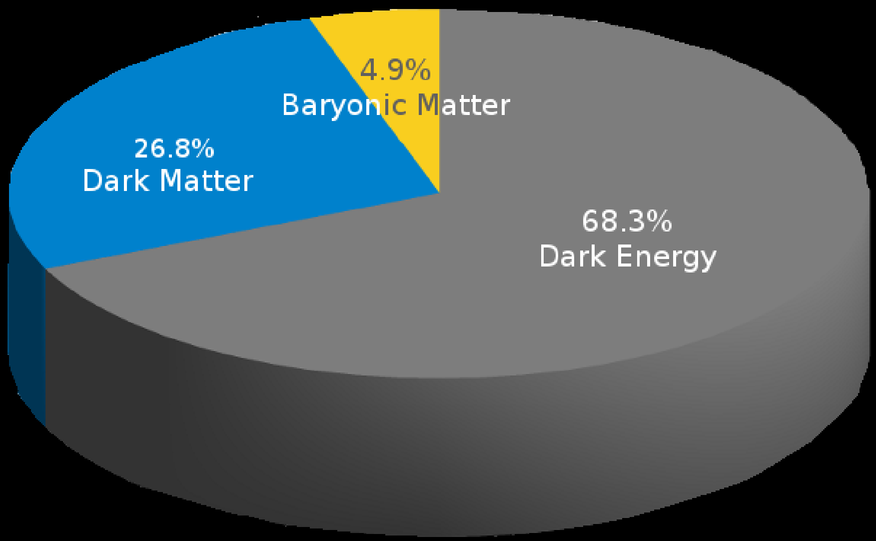
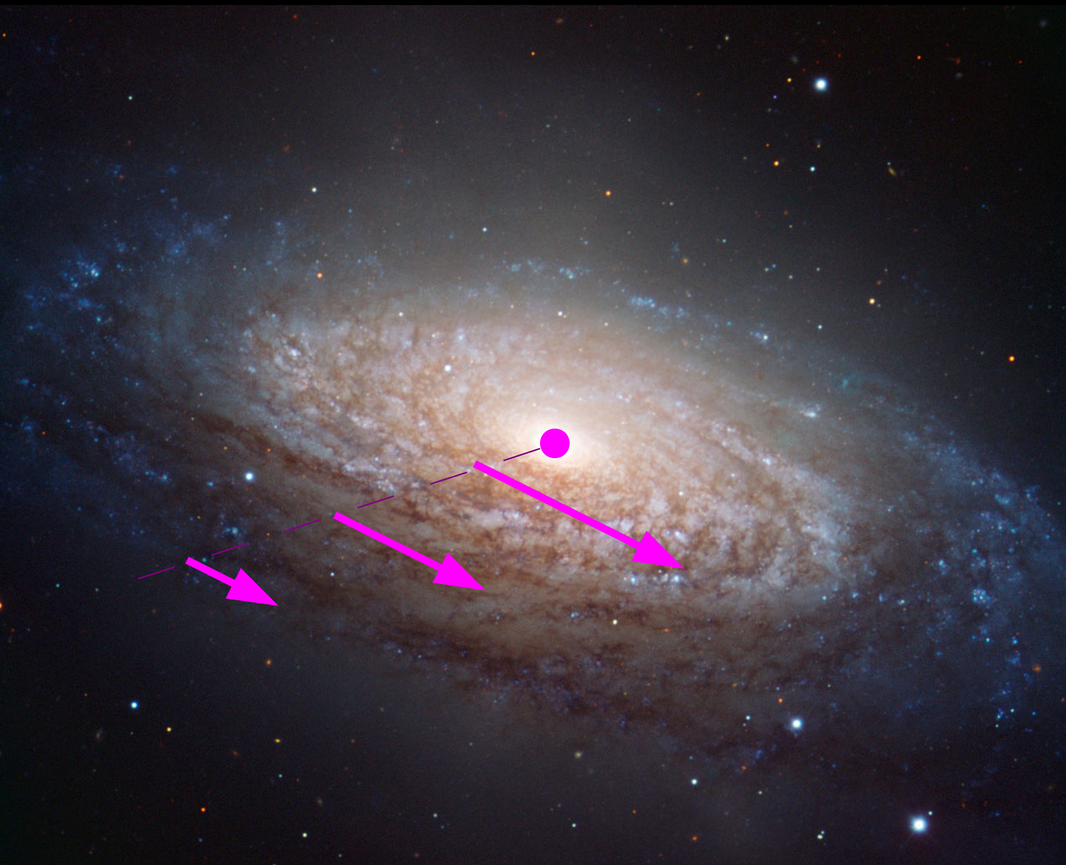


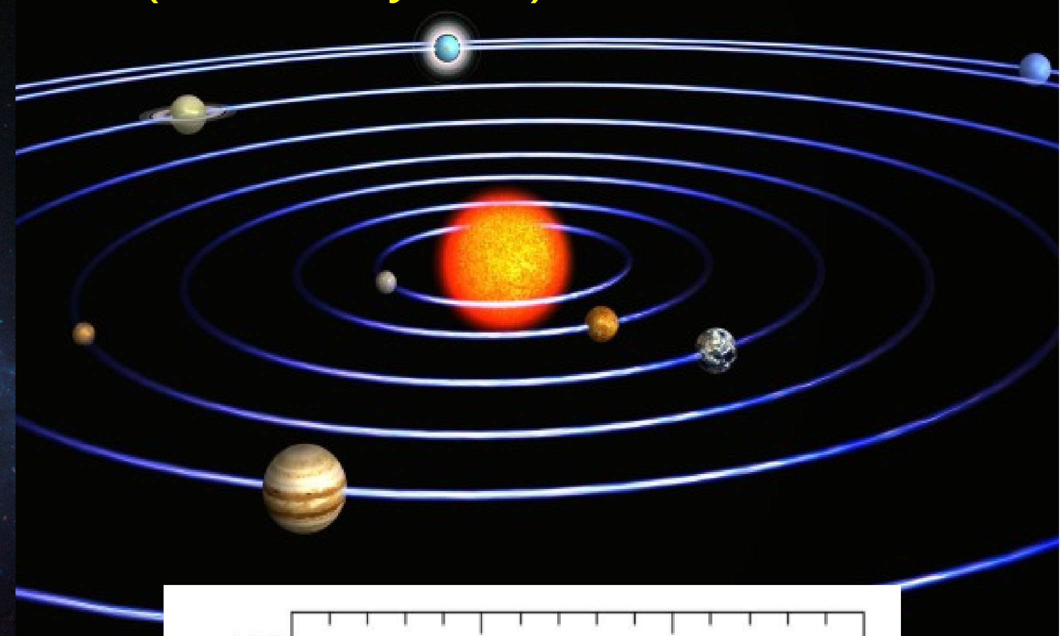
about 100'00 dark matter particles cross an area of 1 cm² per second

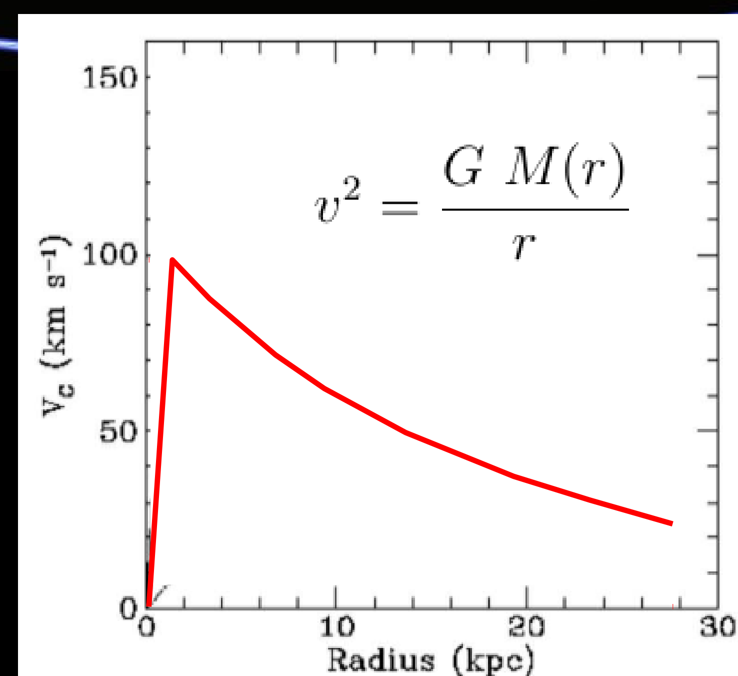

Part 1 –
Evidence for Dark Matter

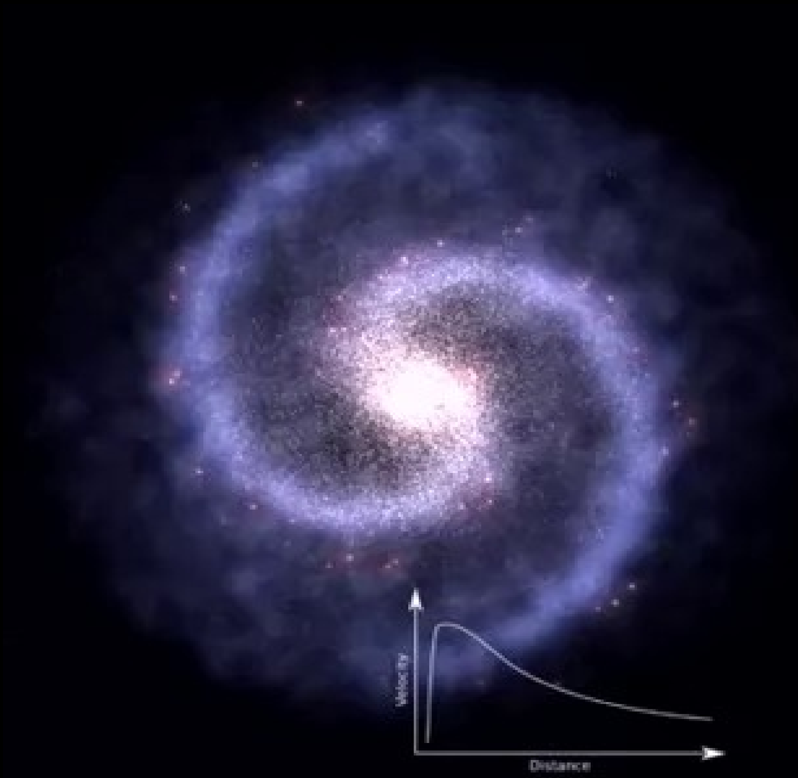


Expect: Kepler Rotation

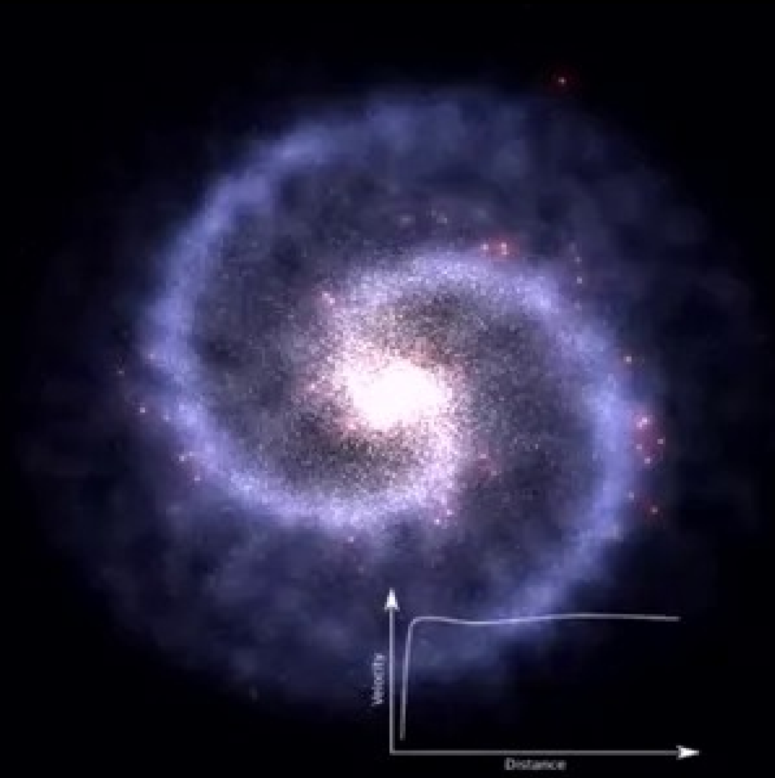
(as solar system)







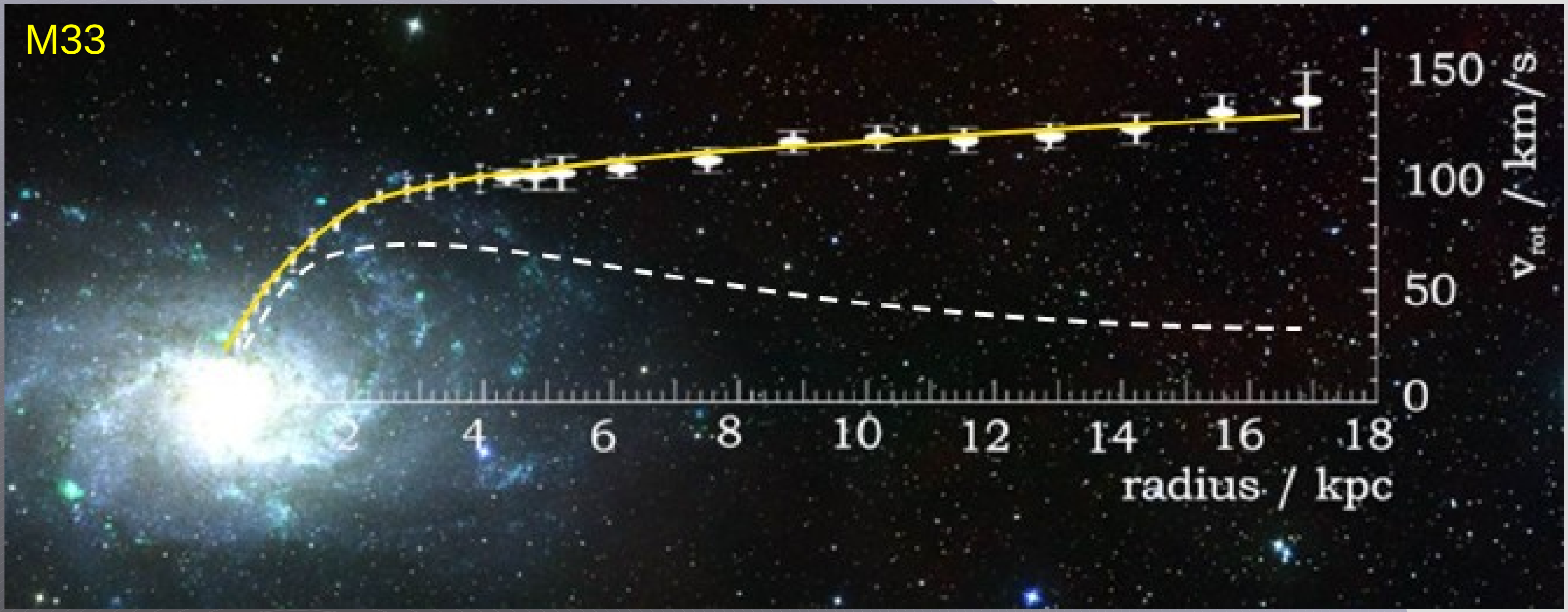
Expectation



Observation

Galactic Rotation Curves

Measurement: flat rotation profile … well beyond visible stars

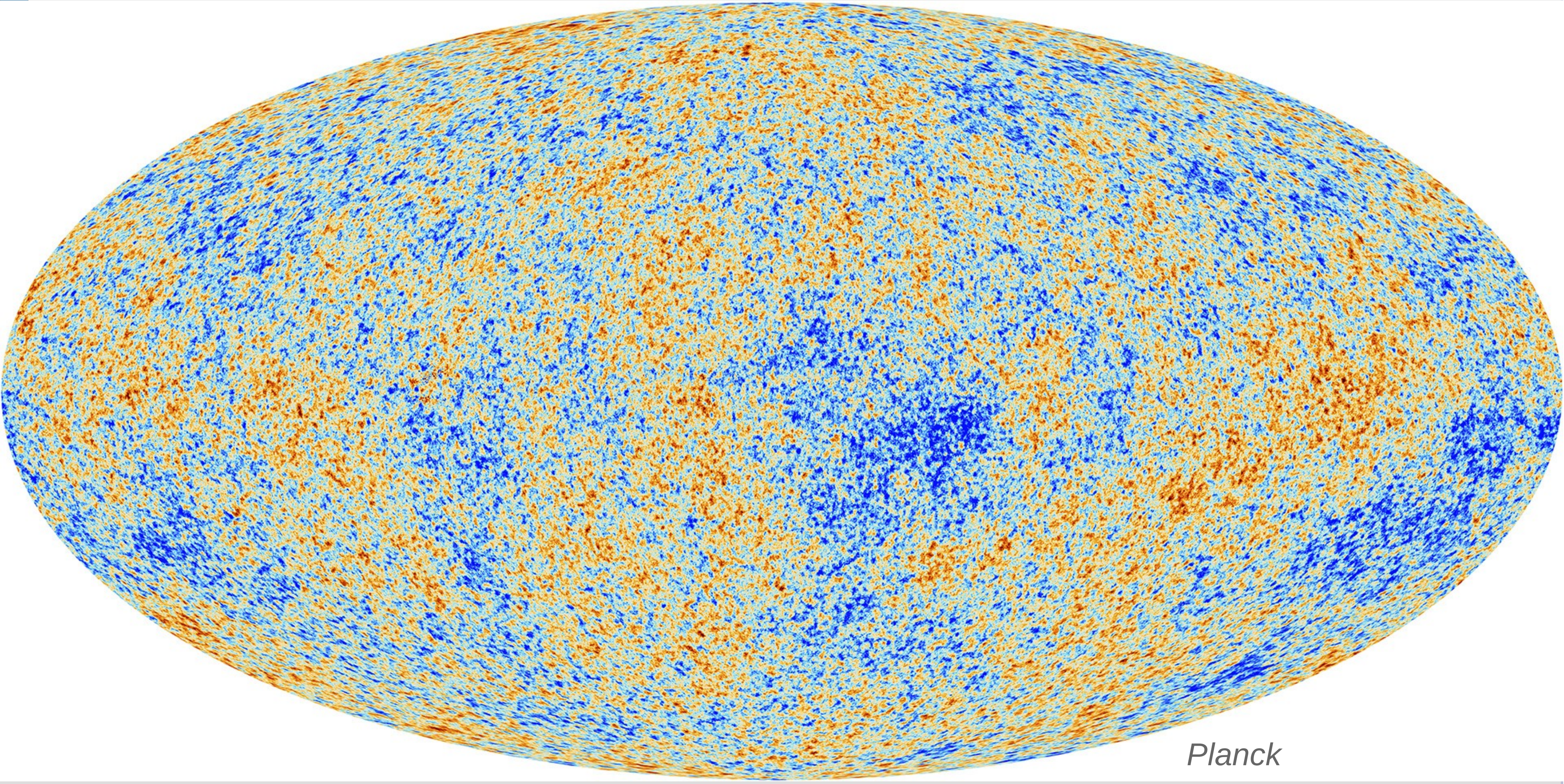


„Halo" made from Dark Matter

(isothermal sphere, $\rho \sim 1/r^2$)

Cosmic Microwave Background

= afterglow of the hot big bang; variations at $\Delta T/T \sim 10^{-5}$ level

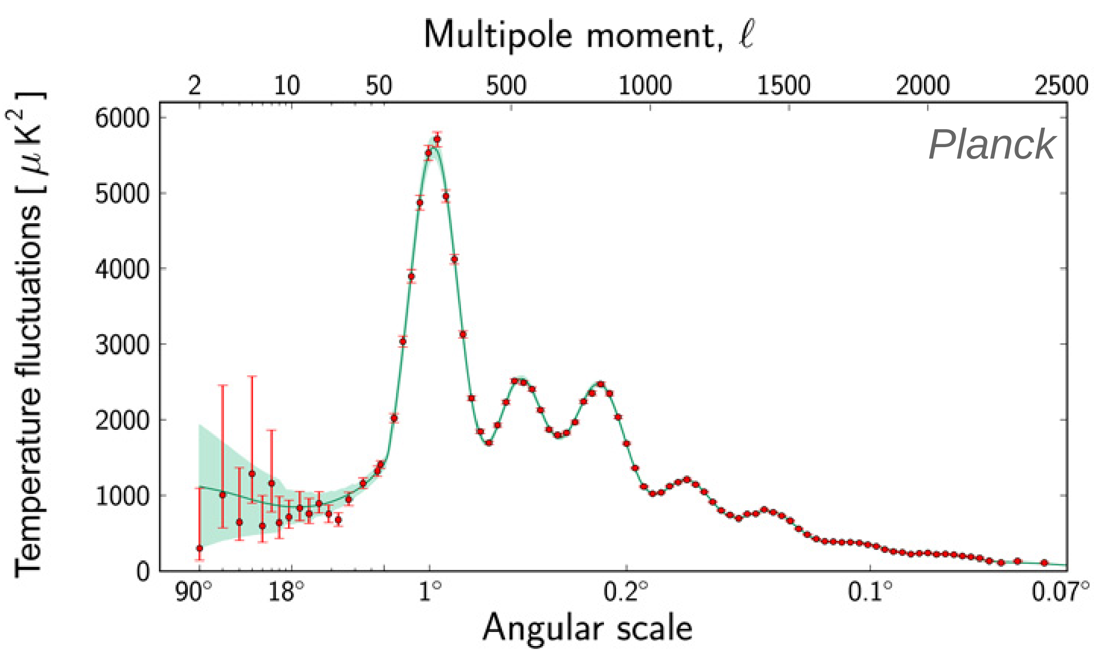


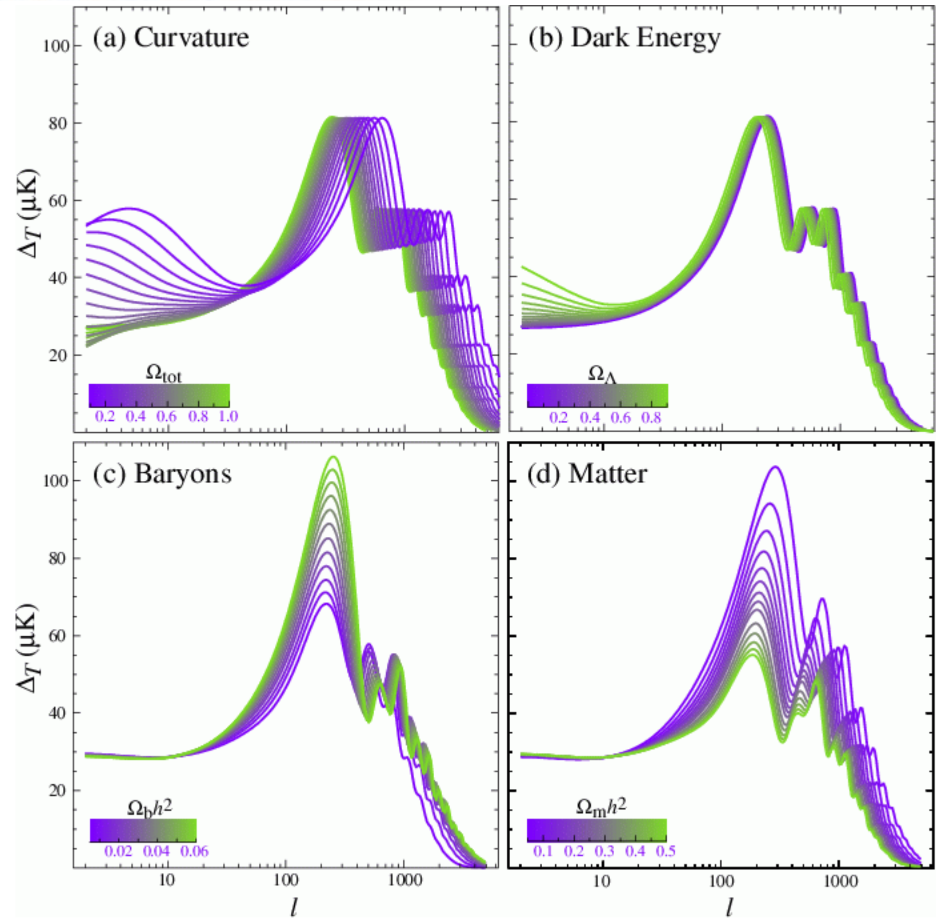
Planck

Cosmic Microwave Background

= afterglow of the hot big bang;   variations at $\Delta T/T \sim 10^{-5}$ level

Correlation Analysis:
„typical *T* variation at typical angular scale“

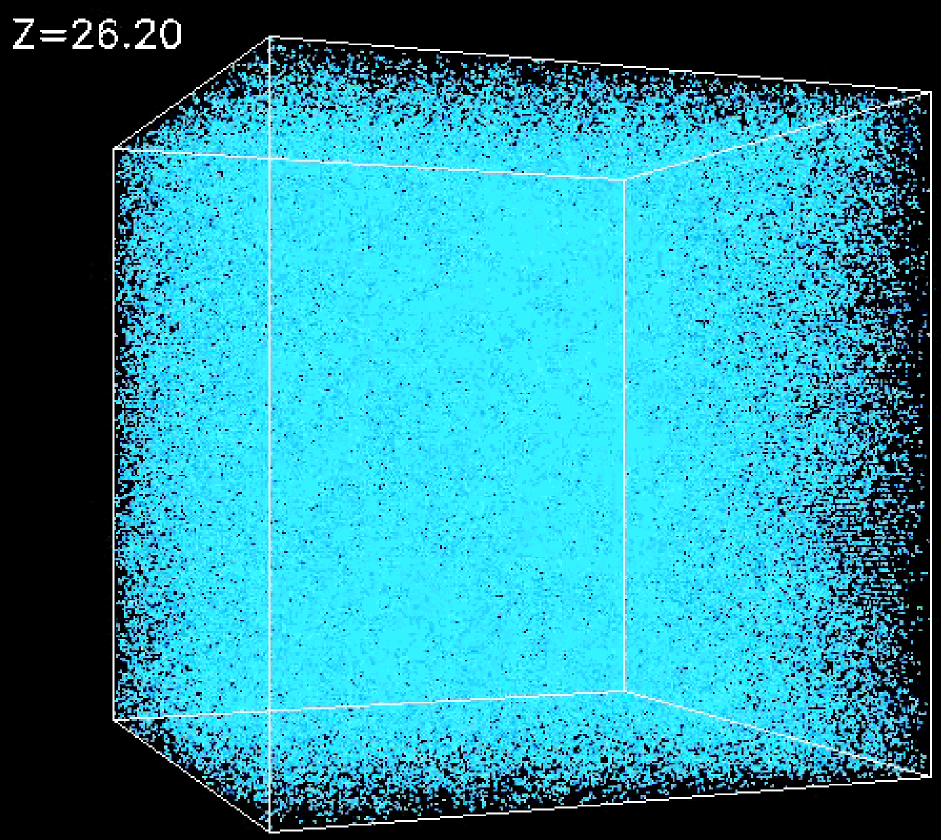


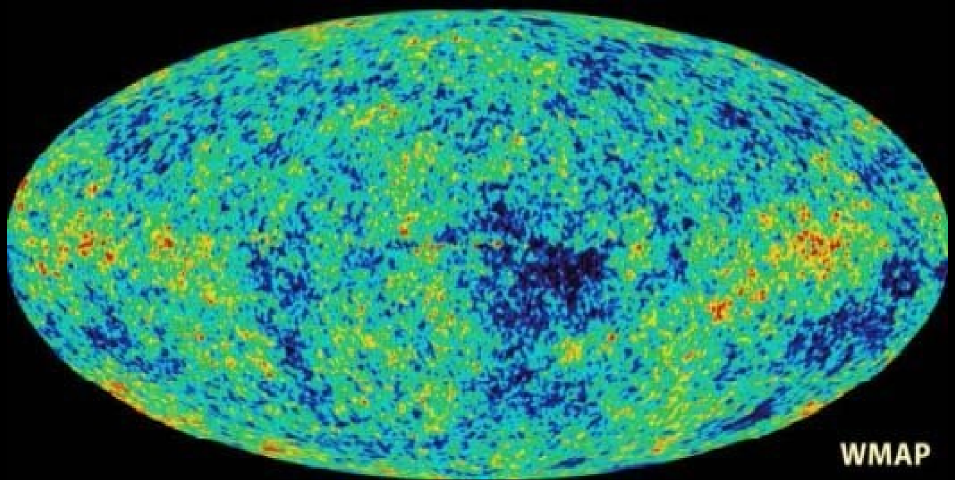


Planck

Dark Matter shapes the Universe

~40M years
after big bang

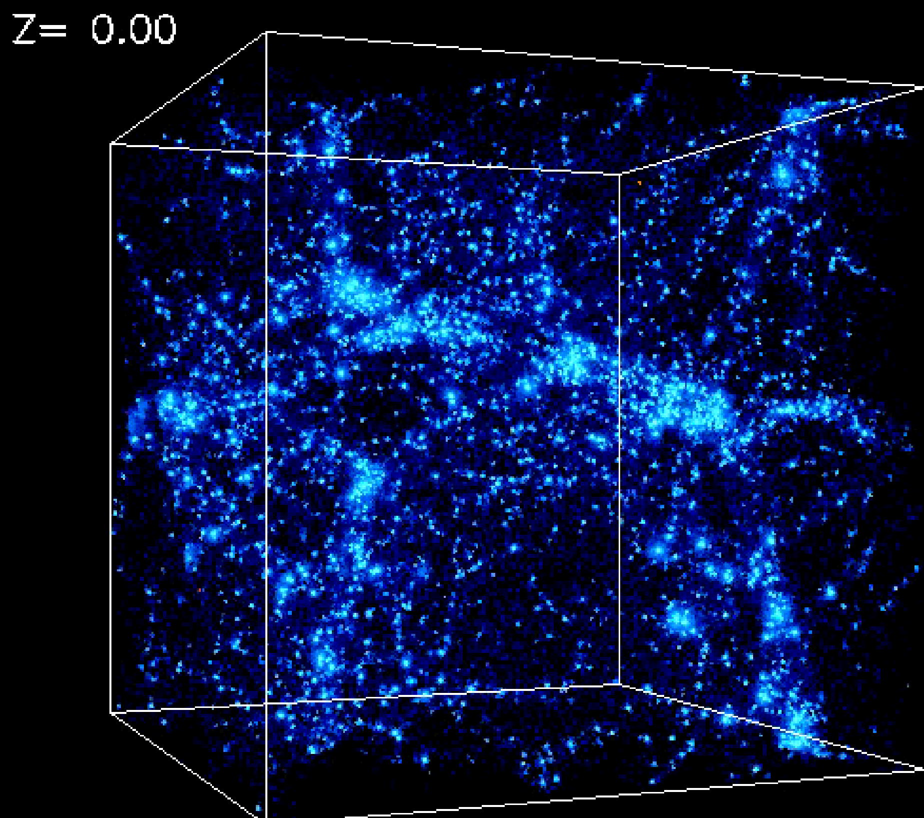


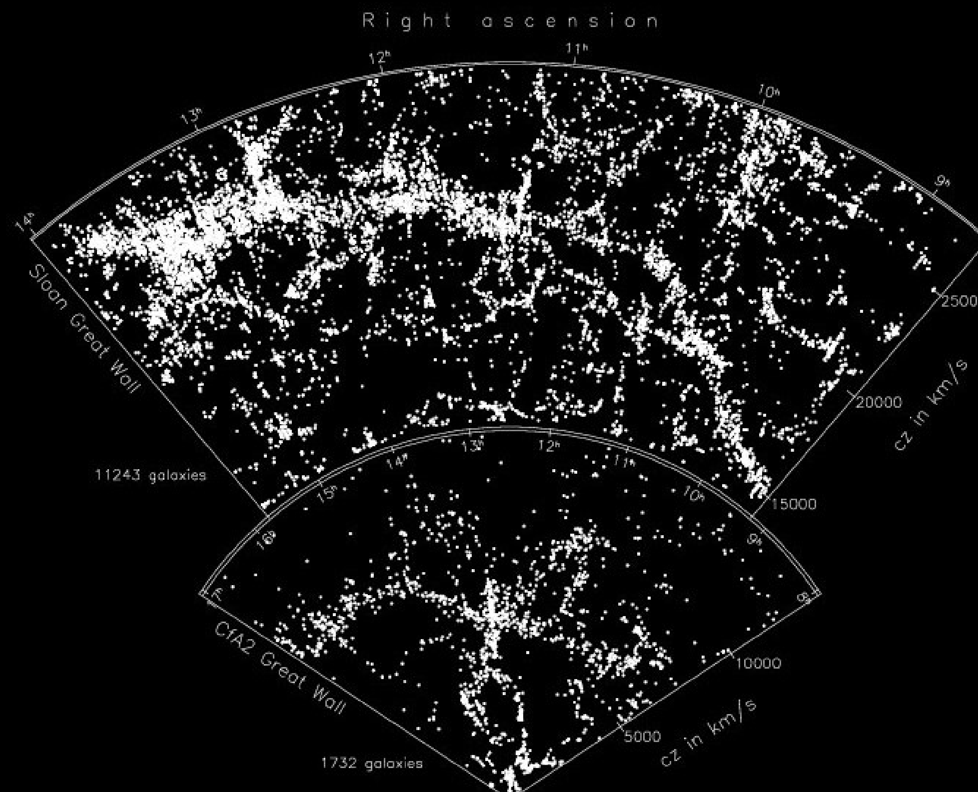


http://cosmicweb.uchicago.edu

Dark Matter shapes the Universe

now





http://cosmicweb.uchicago.edu

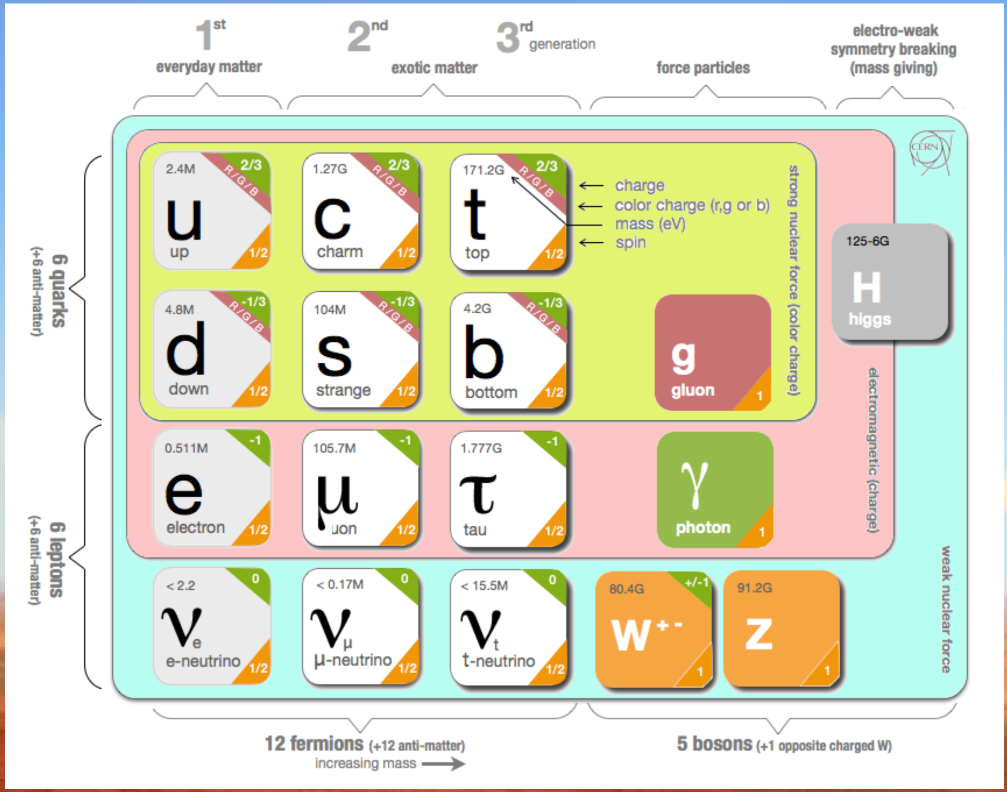
Simulation

Observation (SDSS)

WANTED
FOR MOVING THE UNIVERSE
DARK MATTER

Looking for matter with the following properties:
- „invisible"
- „cold" (= „slow")
- almost collisionless
- stable

REWARD: NOBLE PRICE?



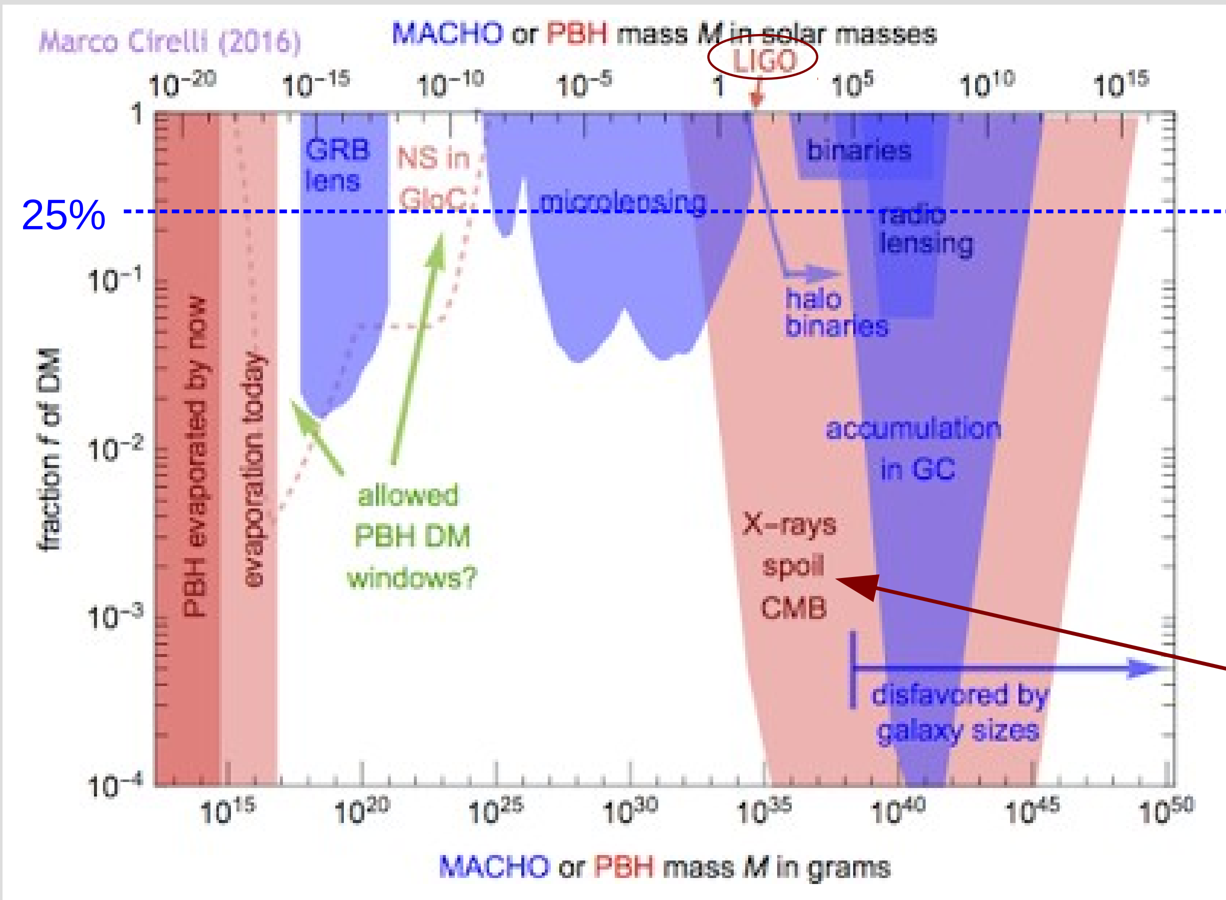
<u>Problem</u>:
no known particle fits the description

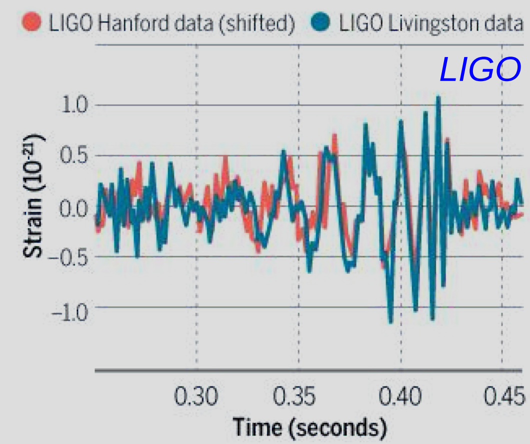
→ **we need to look for something new**

weakly **i**nteracting **m**assive **p**article (WIMP)

Primordial Black Holes?

Can primordial black holes (PBH) formed in the big bang be the dark matter?





Black holes moving in early dense universe accrete matter and produce X-rays
→ ionize atoms
→ affect CMB

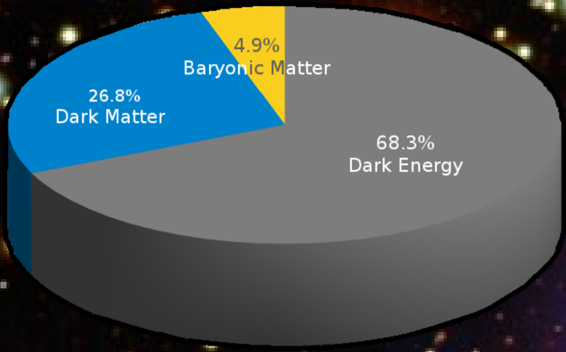
constraints in 10-100 M_{sun} range (LIGO):

– **PBHs cannot constitute >0.01% of dark matter** — *Astrophys.J. 680, 829 (2008)*

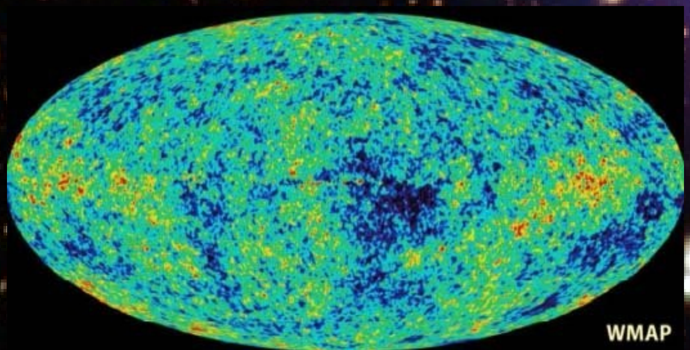
– *but*: new discussion about PBH dark matter started — *PRL 116, 201301 (2016)*

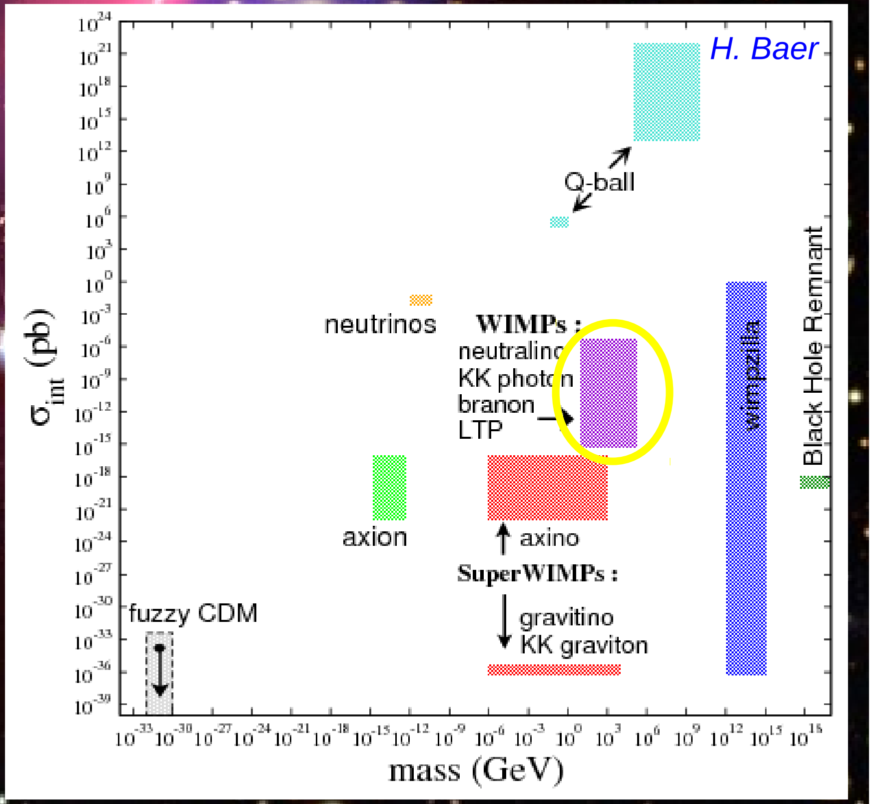
maybe PBH not dark matter but faster merger rate — *PRL 117, 061101 (2016)*

Dark Matter: (indirect) Evidence



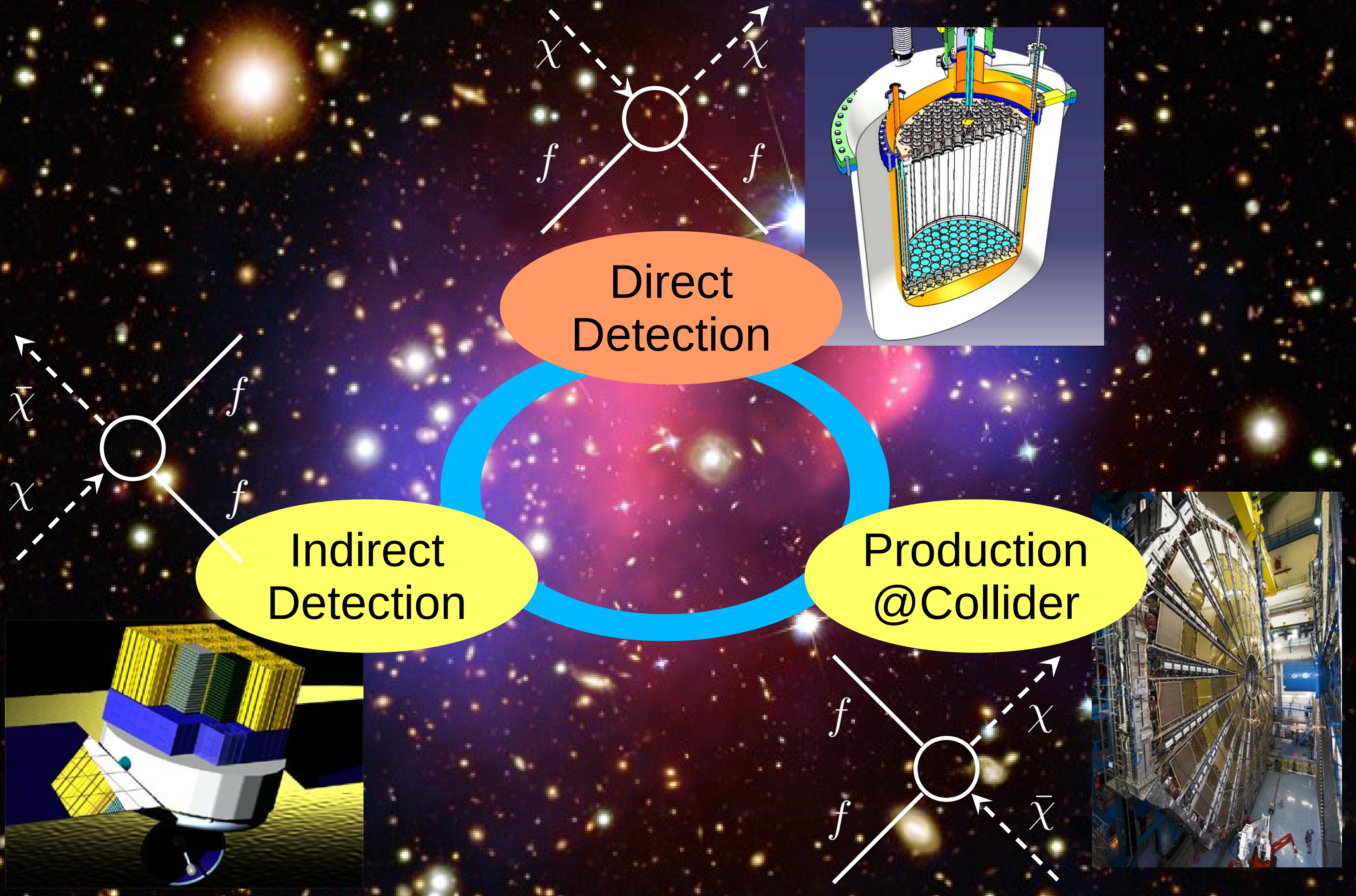
Particle Dark Matter Candidates:
- **WIMP** → *„WIMP miracle“*
- Axion
- SuperWIMPs
- sterile neutrinos
- WIMPless dark matter
- Gravitino
- ...





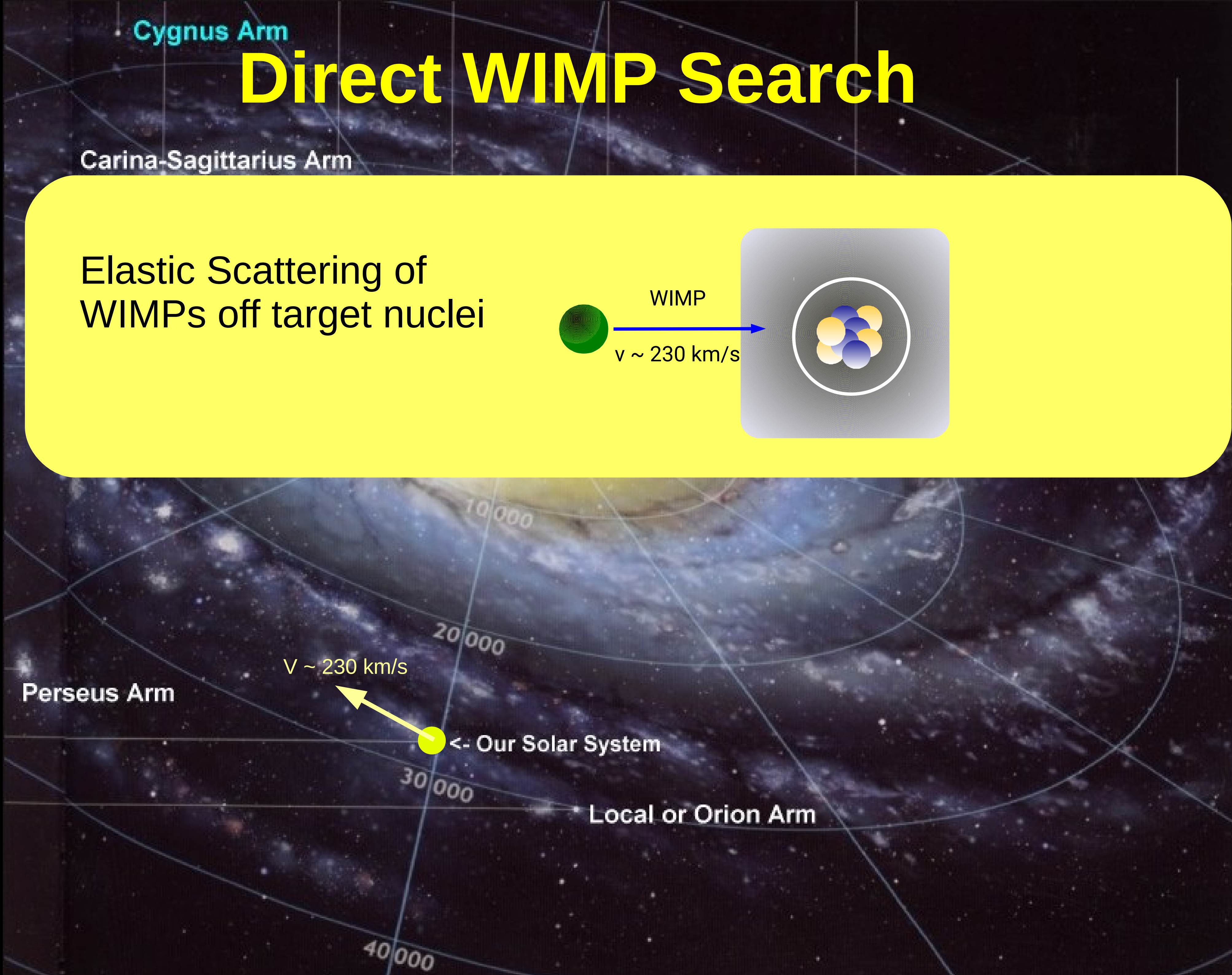

Part 2 –
Searching for Dark Matter

Dark Matter Search



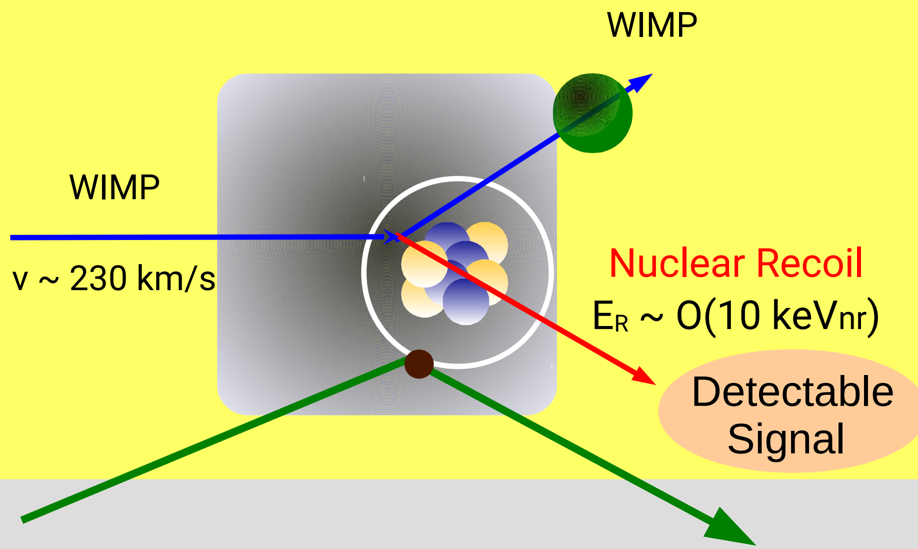
Direct WIMP Search

Elastic Scattering of WIMPs off target nuclei



Direct WIMP Search

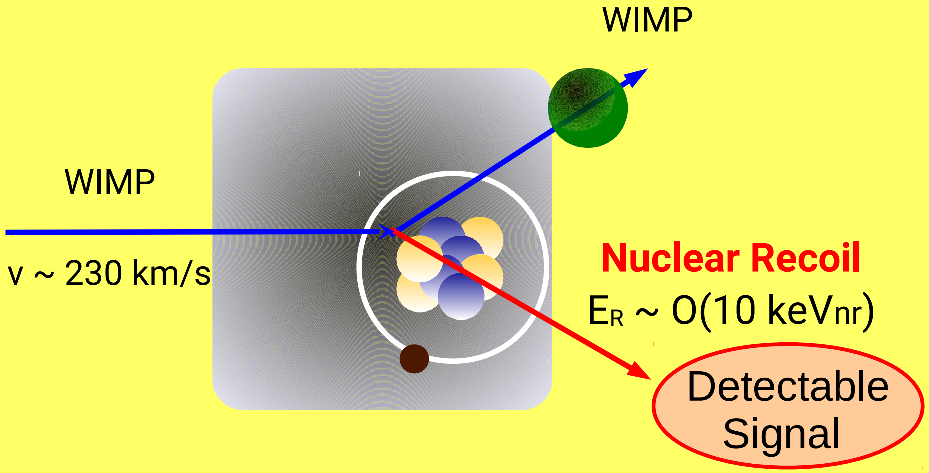
Elastic Scattering of WIMPs off target nuclei
➡ nuclear recoil

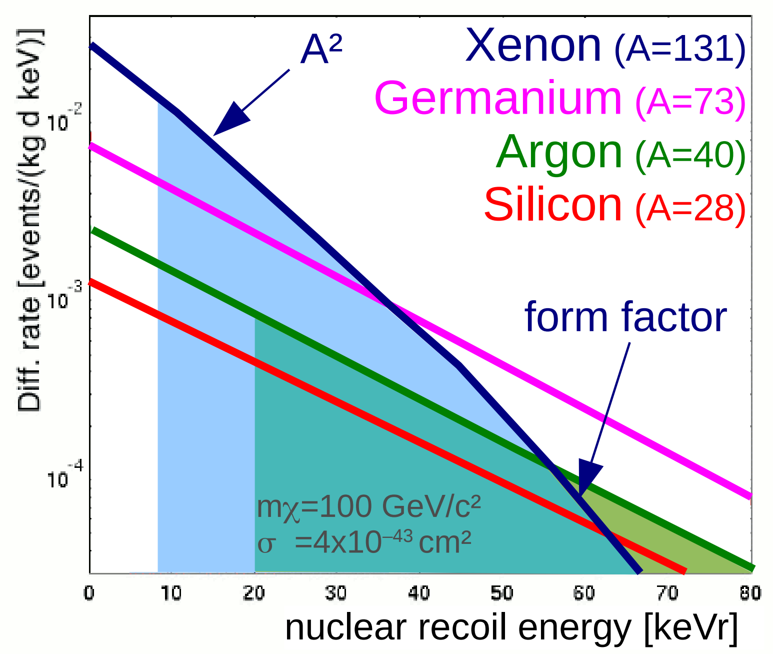


gamma- and beta-particles (background) interact with the atomic electrons
→ **electronic recoil** [in keVee]

Direct WIMP Search

Elastic Scattering of WIMPs off target nuclei
➡ nuclear recoil



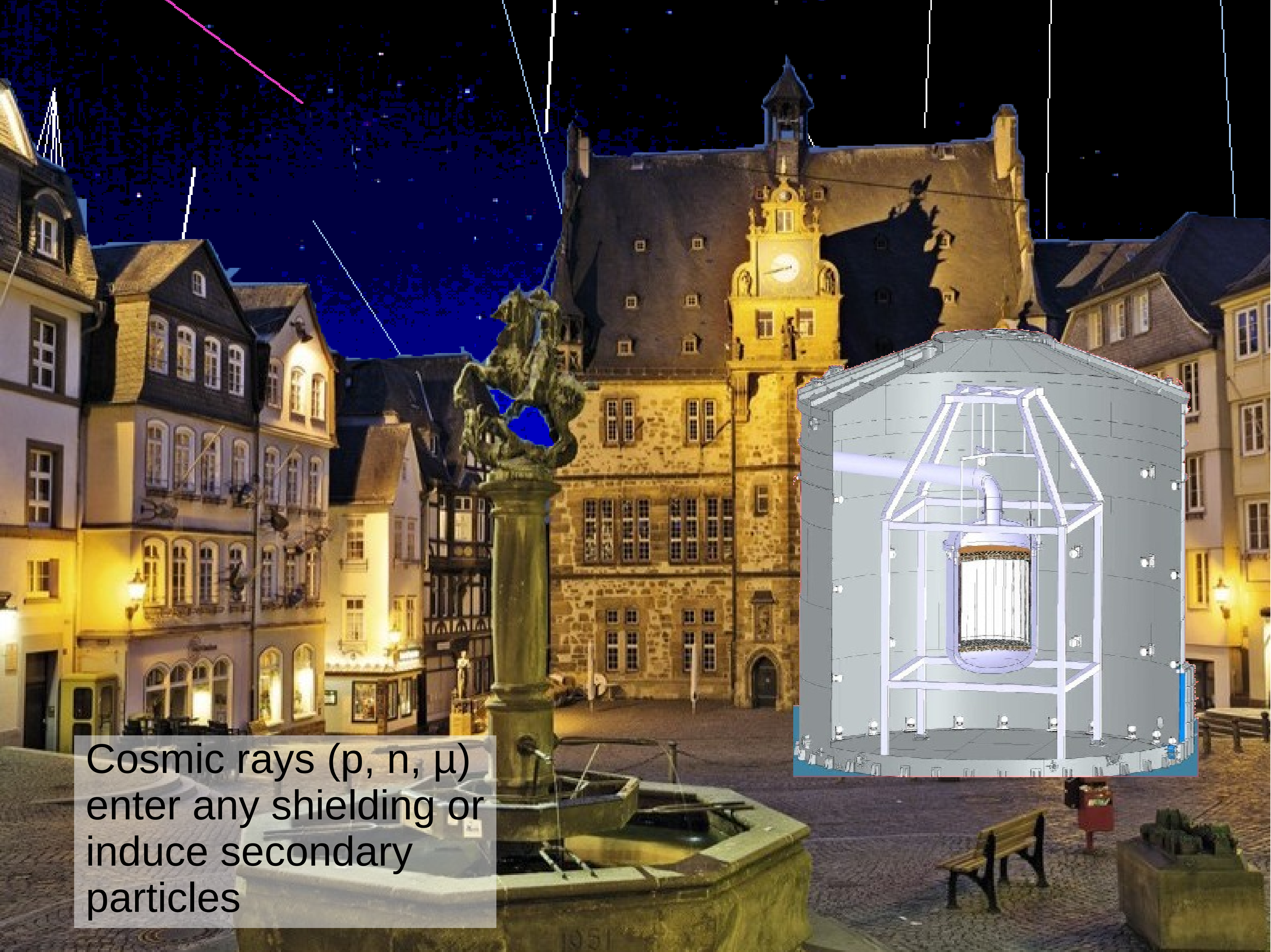


Event rate $R \propto N \frac{\rho_\chi}{m_\chi} \langle \sigma_{\chi-N} \rangle$

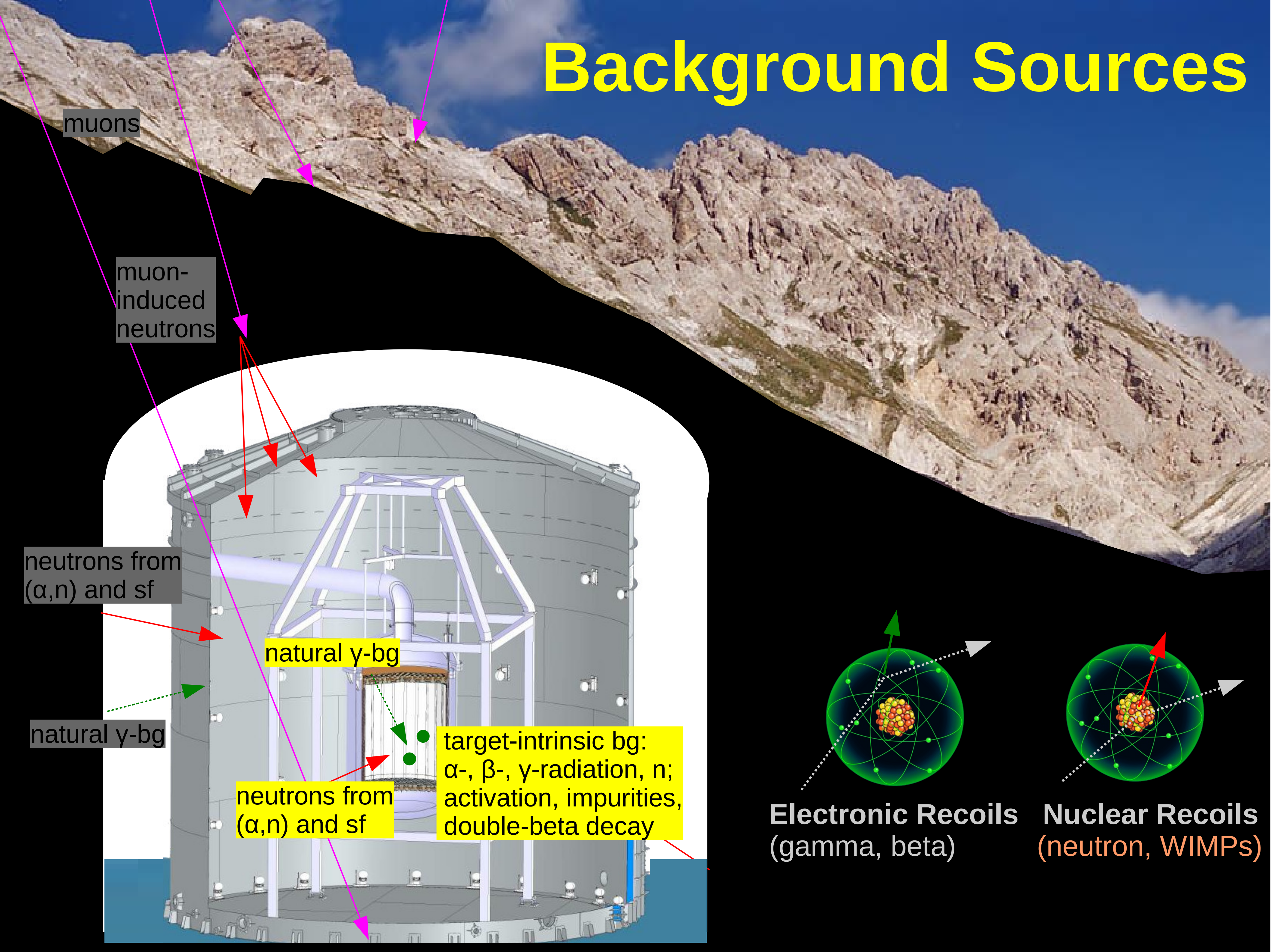
Detector — Local DM Density — Physics

$\rho_\chi \sim 0.3$ GeV/c²

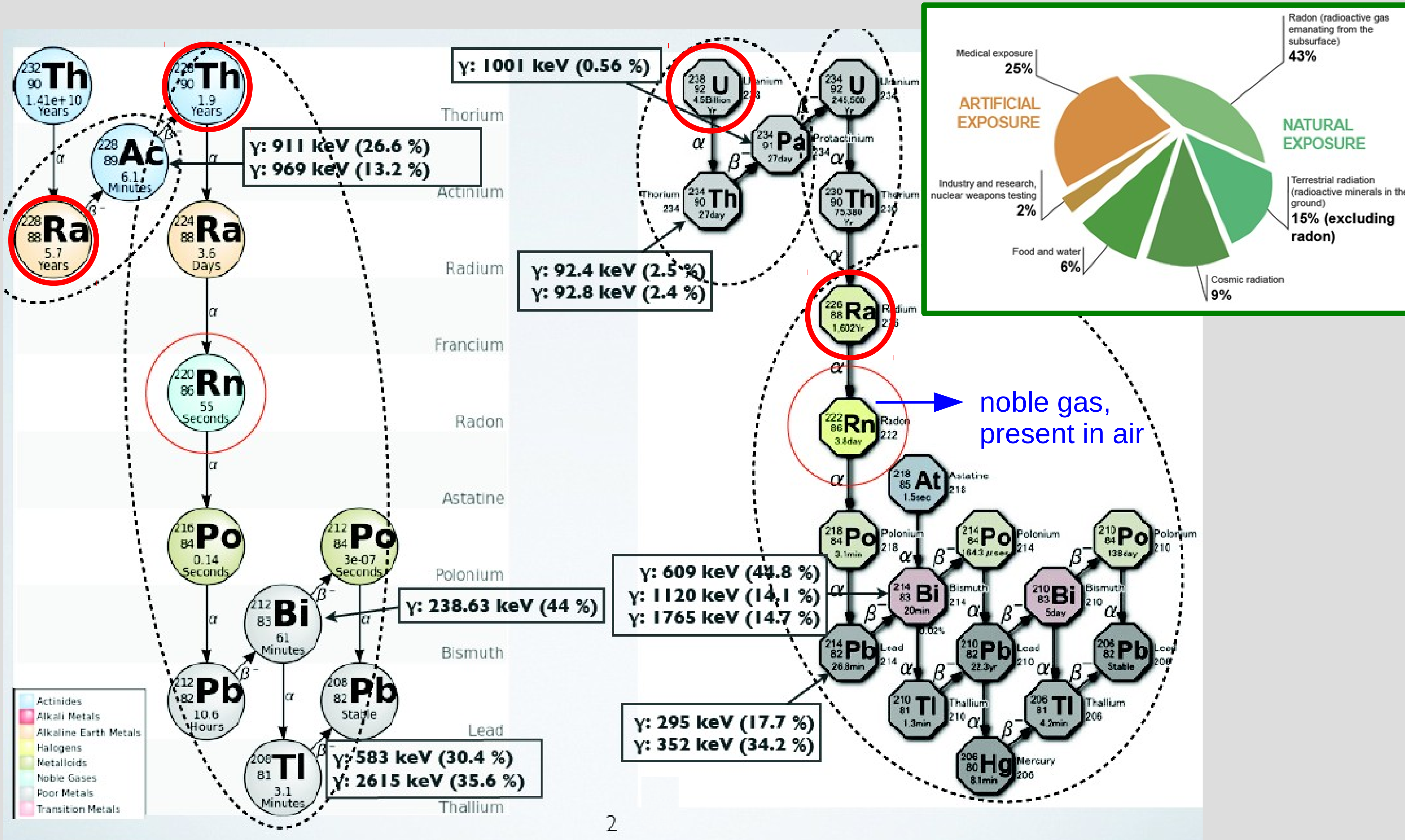
→ very small: ≪1 event/kg/year
→ search for rare events
→ **low-background crucial**


Cosmic rays (p, n, μ)
enter any shielding or
induce secondary
particles

Background Sources



The U and Th Chains

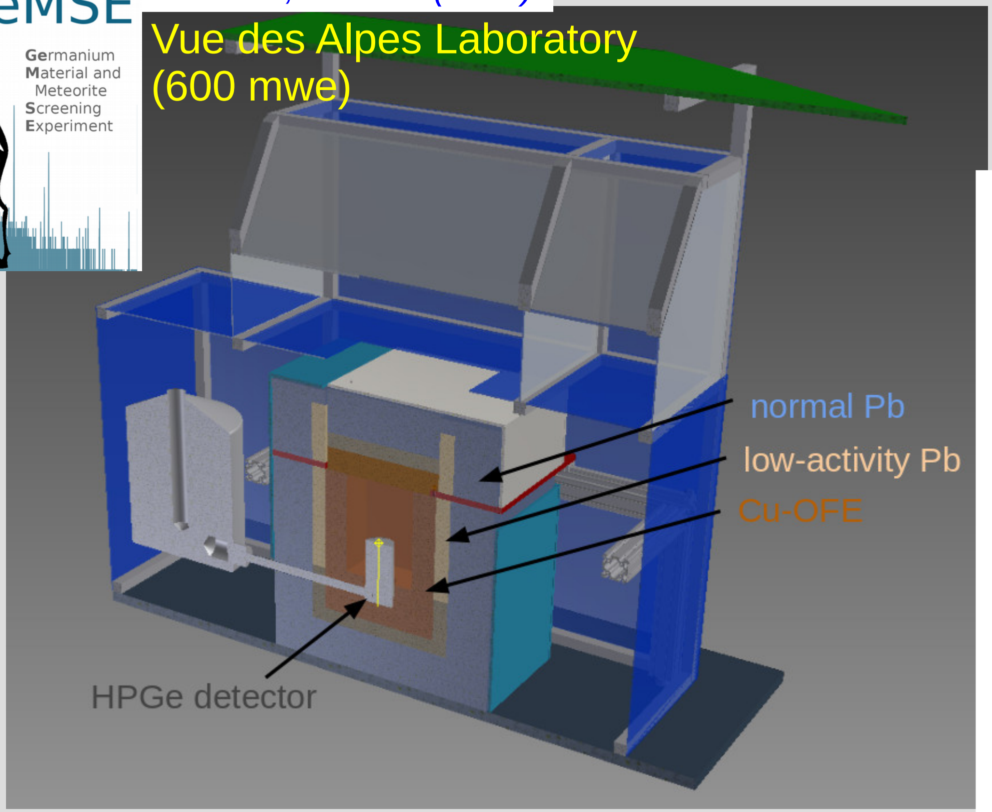


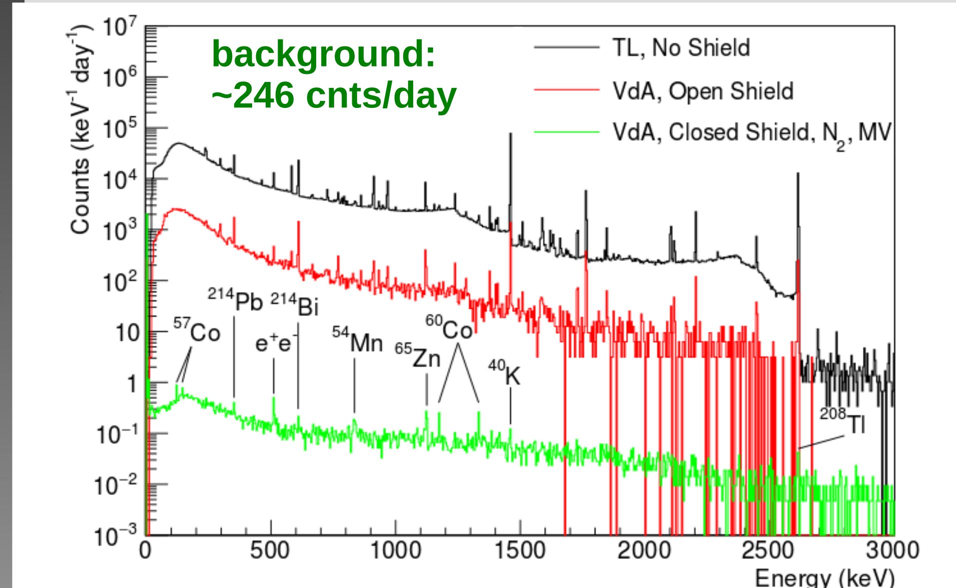

GeMSE
Germanium
Material and
Meteorite
Screening
Experiment

Low-background Screening

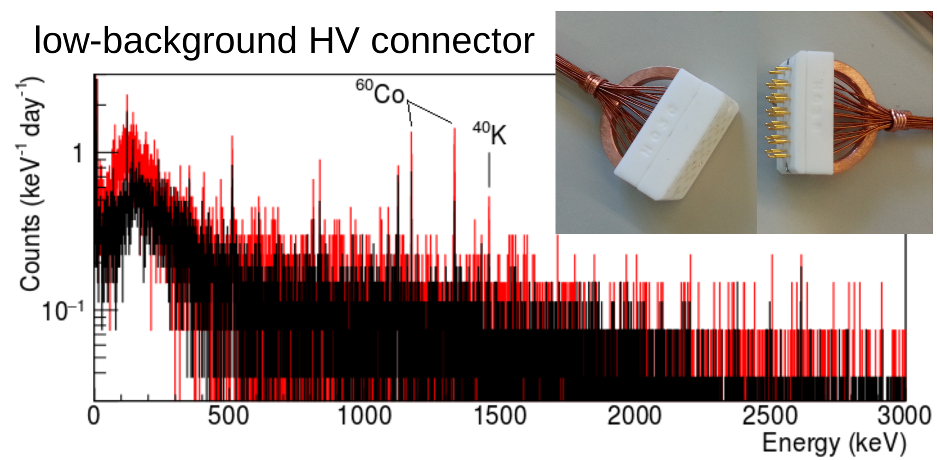
GeMSE
Germanium Material and Meteorite Screening Experiment

JINST 11, P12017 (2016)



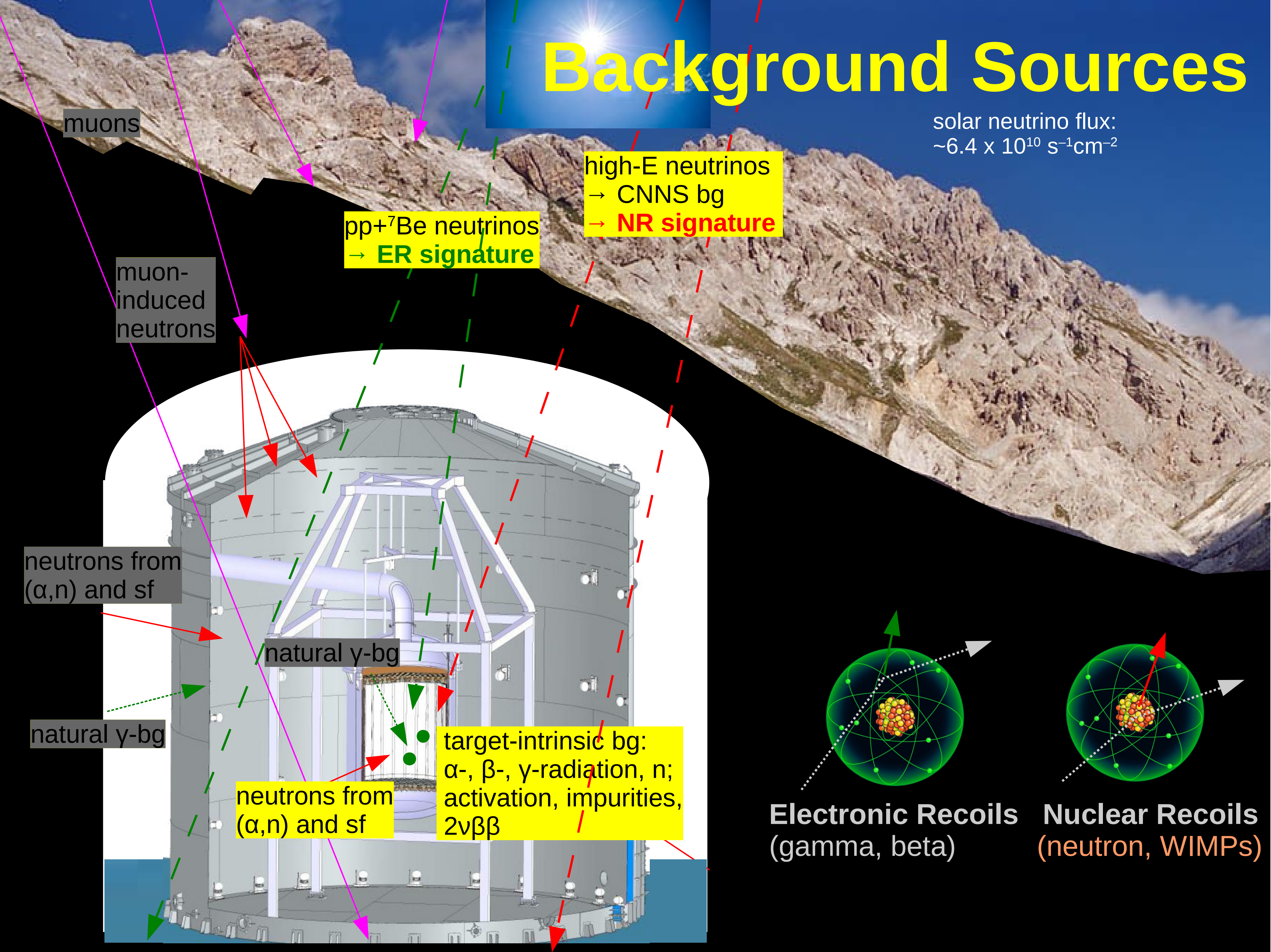


low-background HV connector



Identify materials with lowest radioactivity:

- γ-spectrometry using HPGe Detectors
- mass spectroscopy: ICP-MS, GDMS
- neutron activation analysis
- ^{222}Rn emanation



Background Sources
solar neutrino flux:
~6.4 x 10^10 s^-1cm^-2
muons
high-E neutrinos
→ CNNS bg
→ NR signature
pp+^7Be neutrinos
→ ER signature
muon-
induced
neutrons
neutrons from
(α,n) and sf
natural γ-bg
natural γ-bg
target-intrinsic bg:
α-, β-, γ-radiation, n;
activation, impurities,
2νββ
neutrons from
(α,n) and sf
Electronic Recoils
(gamma, beta)
Nuclear Recoils
(neutron, WIMPs)

Background Suppression

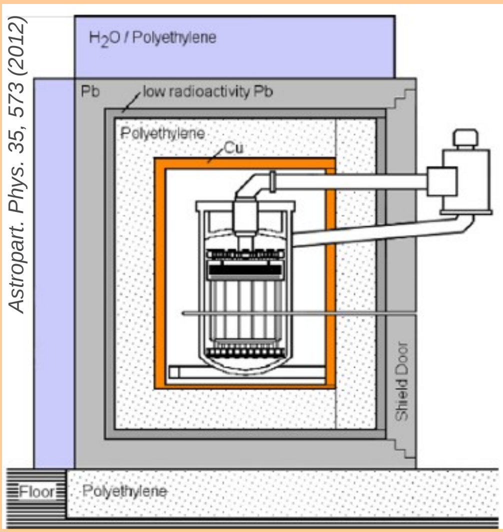
Avoid Backgrounds

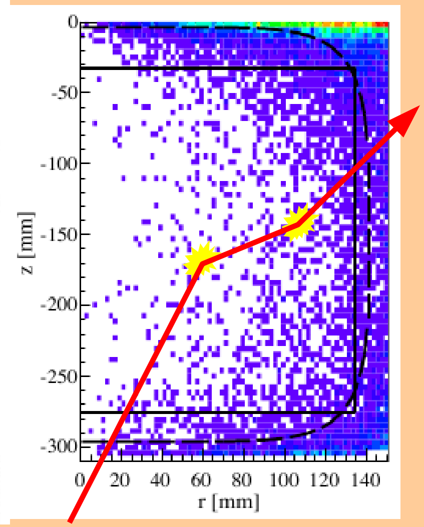
Shielding

deep underground location
large shield (Pb, water, PE)
active veto (μ, γ coincidence)

self shielding → fiducialization

Use of radiopure materials





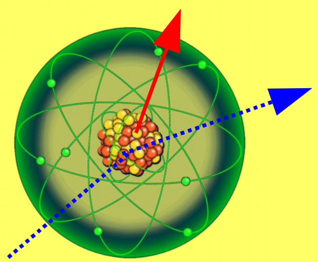
Use knowledge about expected WIMP signal

WIMPs interact only once

→ single scatter selection
requires some position resolution

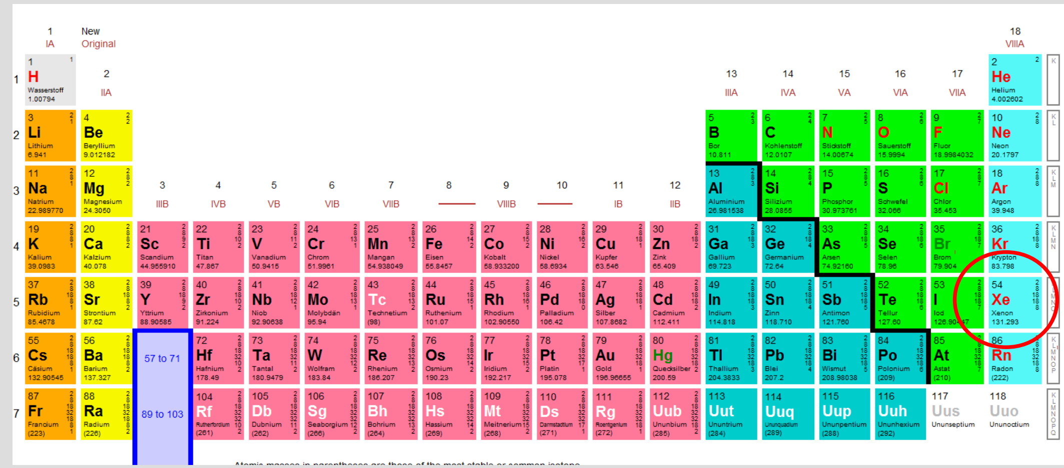
WIMPs interact with target nuclei

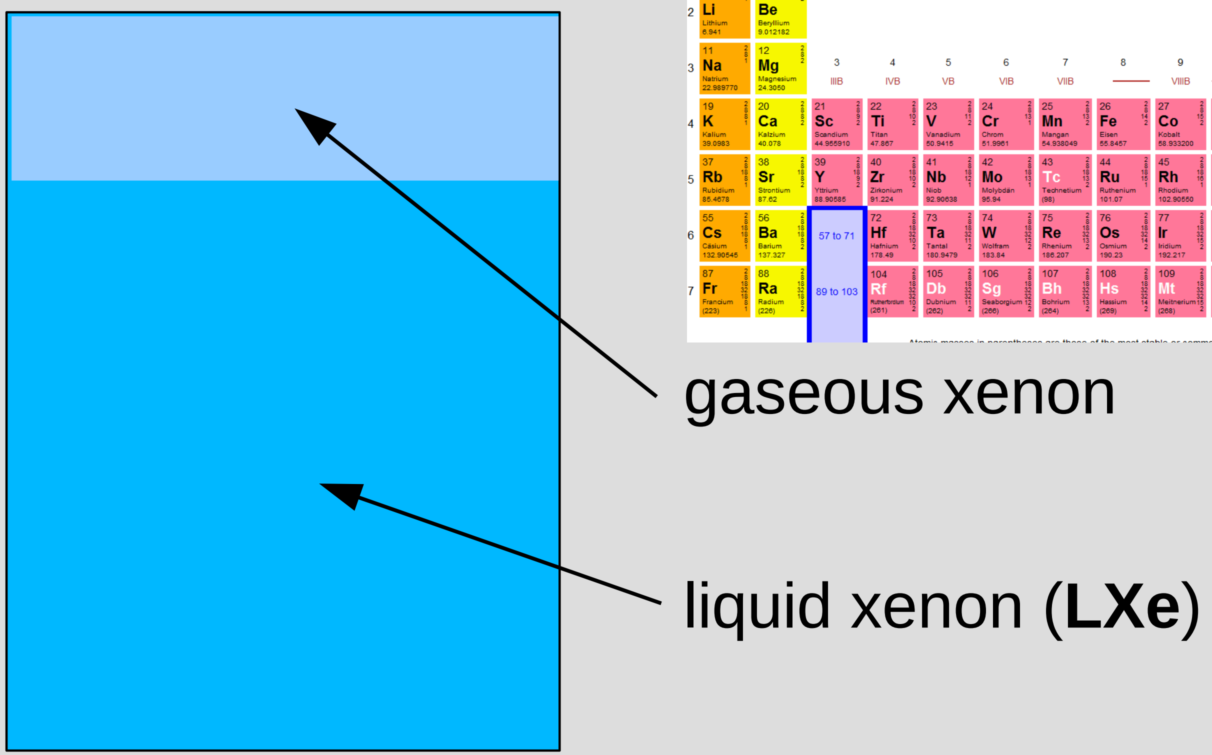
→ nuclear recoils
exploit different d*E*/d*x* from
signal and background



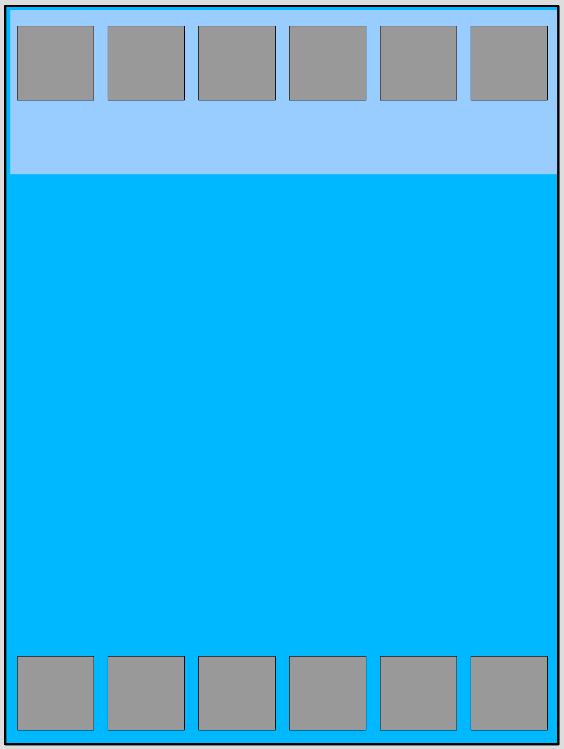
Part 3 – The XENON1T Experiment

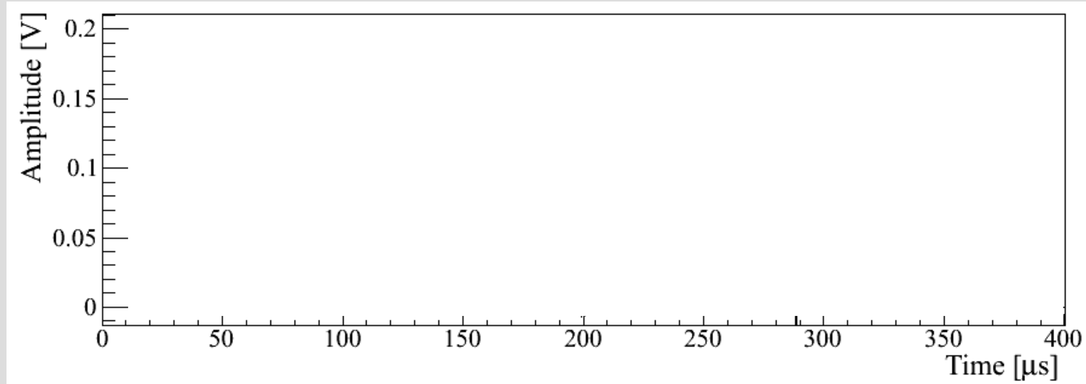
Dual Phase **liquid xenon** TPC

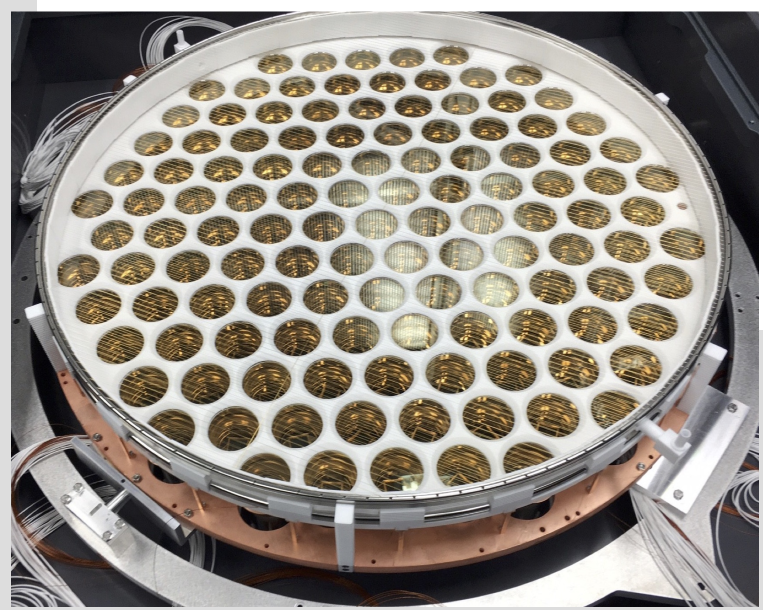




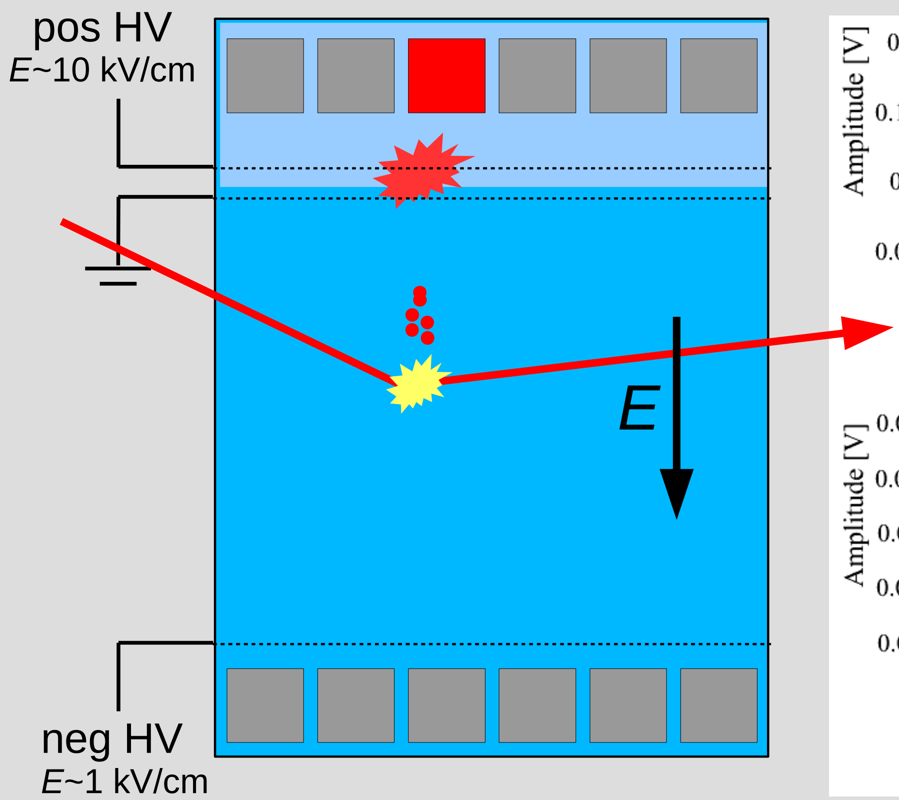
Dual Phase TPC

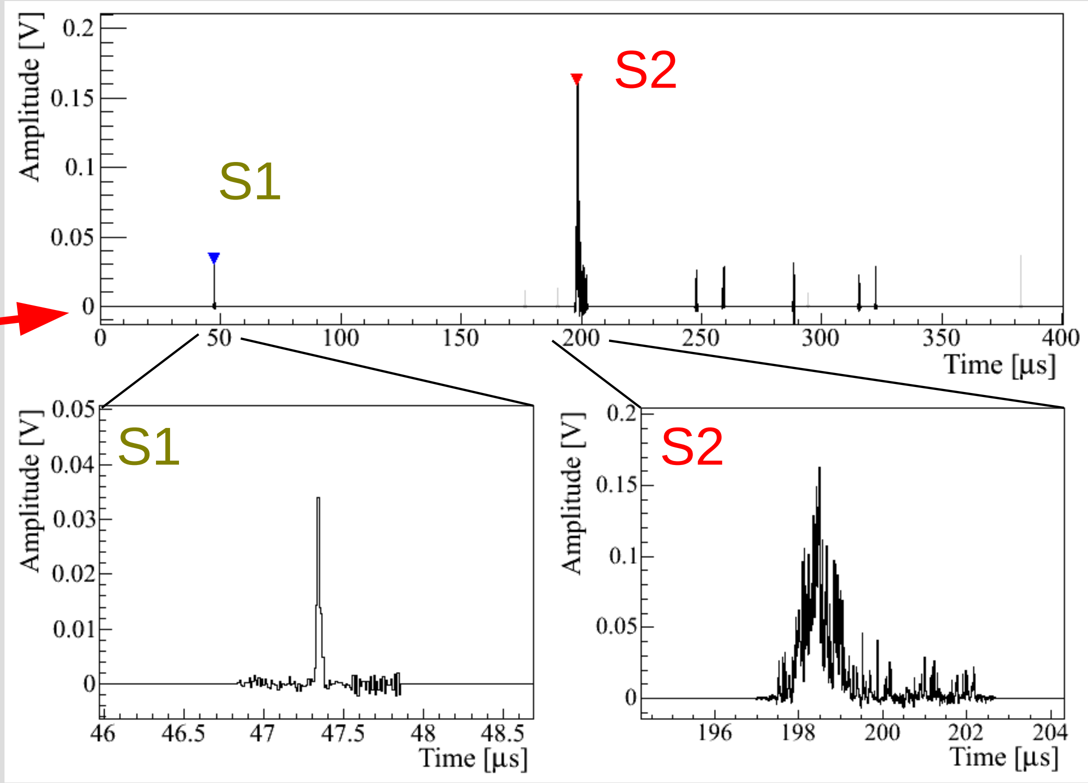




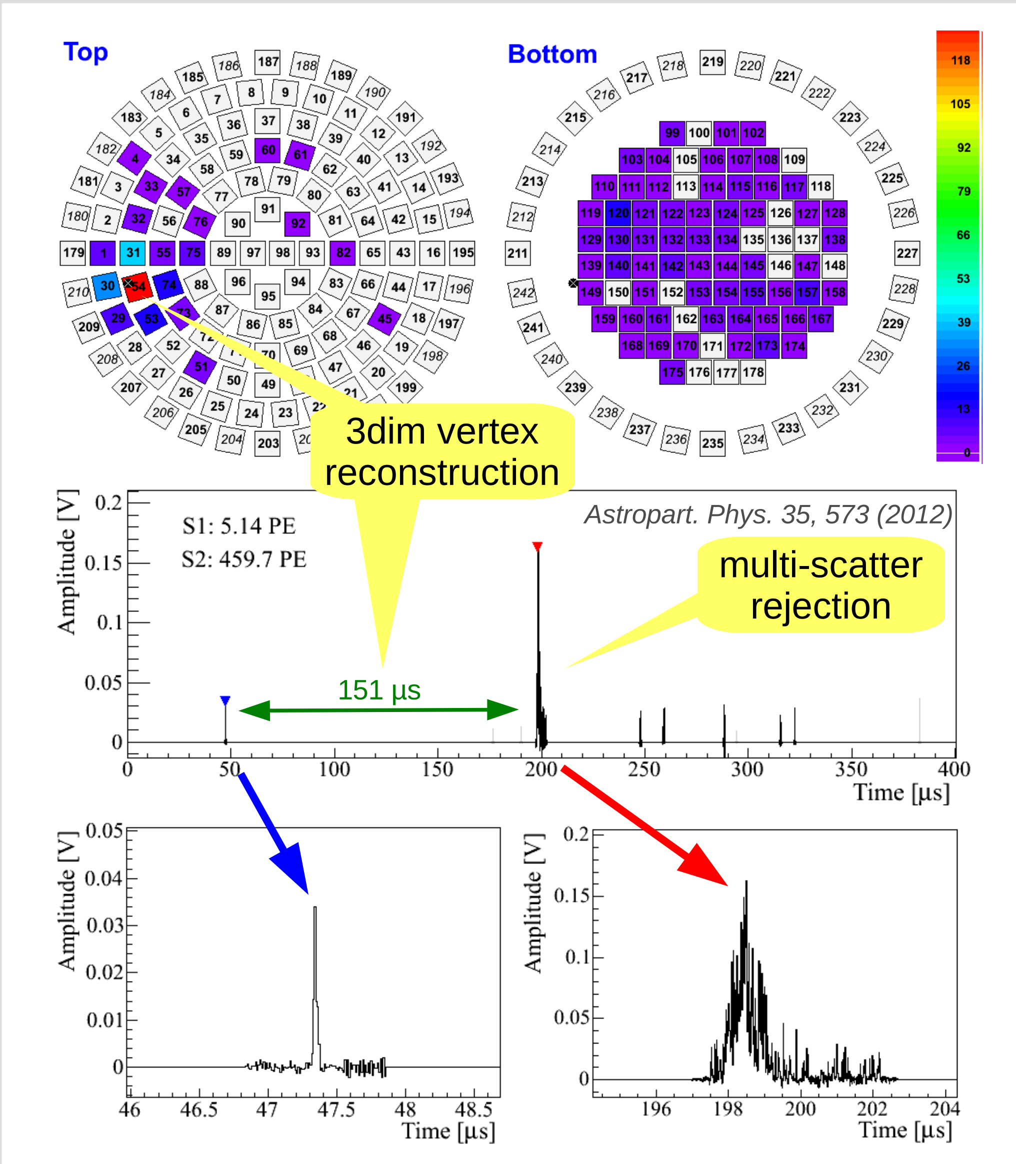


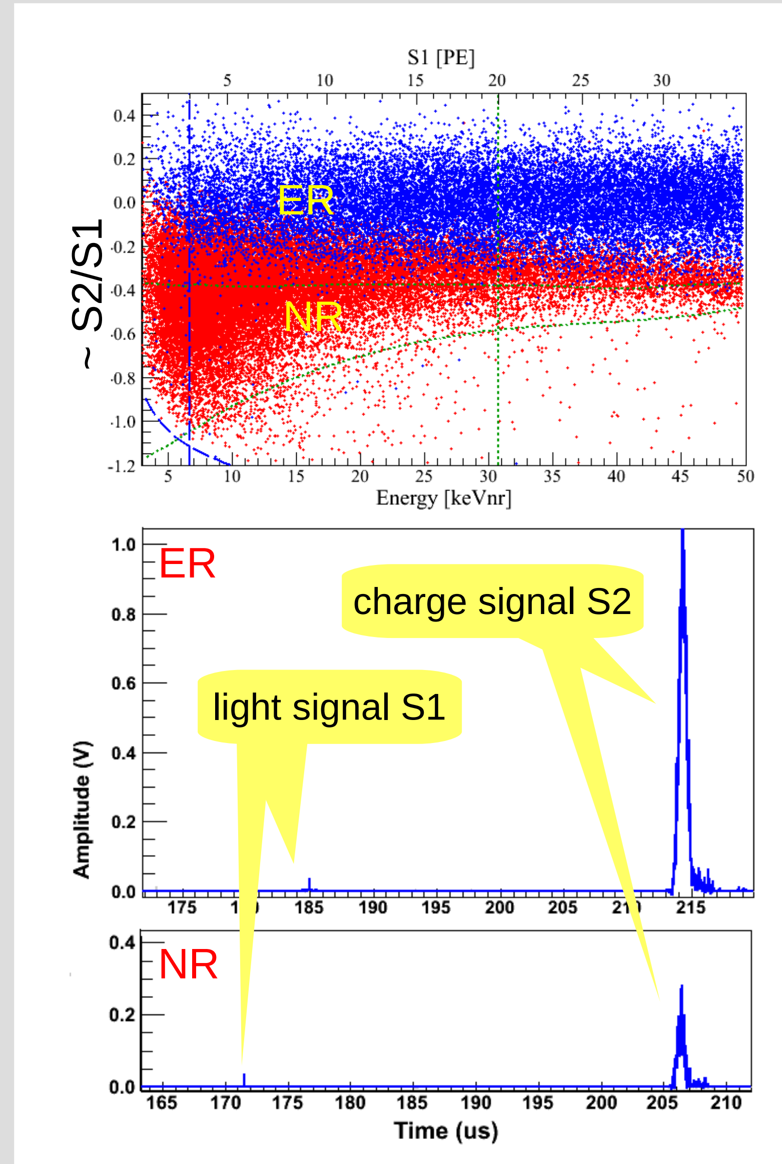
Dual Phase TPC





Dual Phase TPC

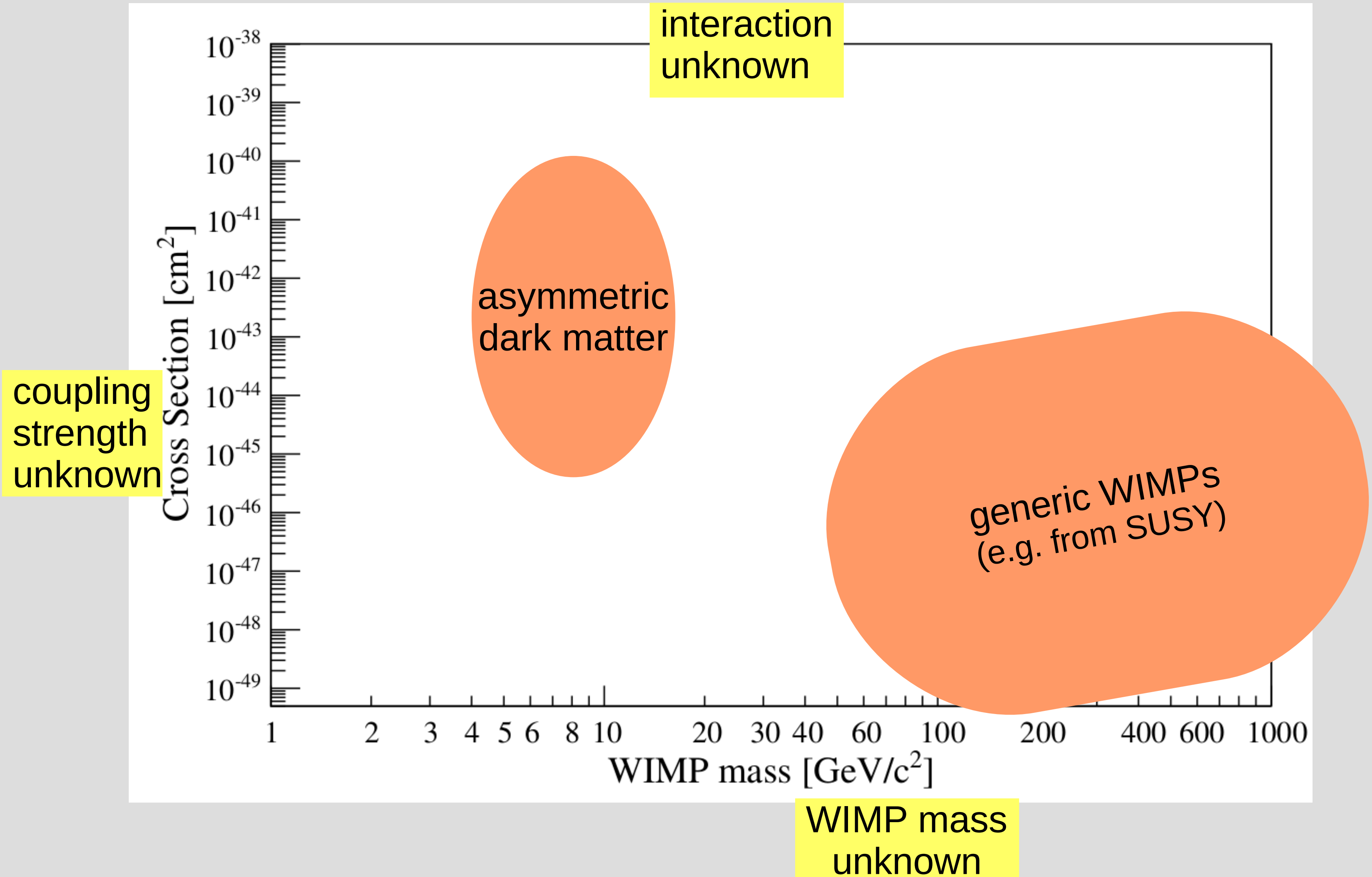




Figures from XENON100

The WIMP Parameter Space

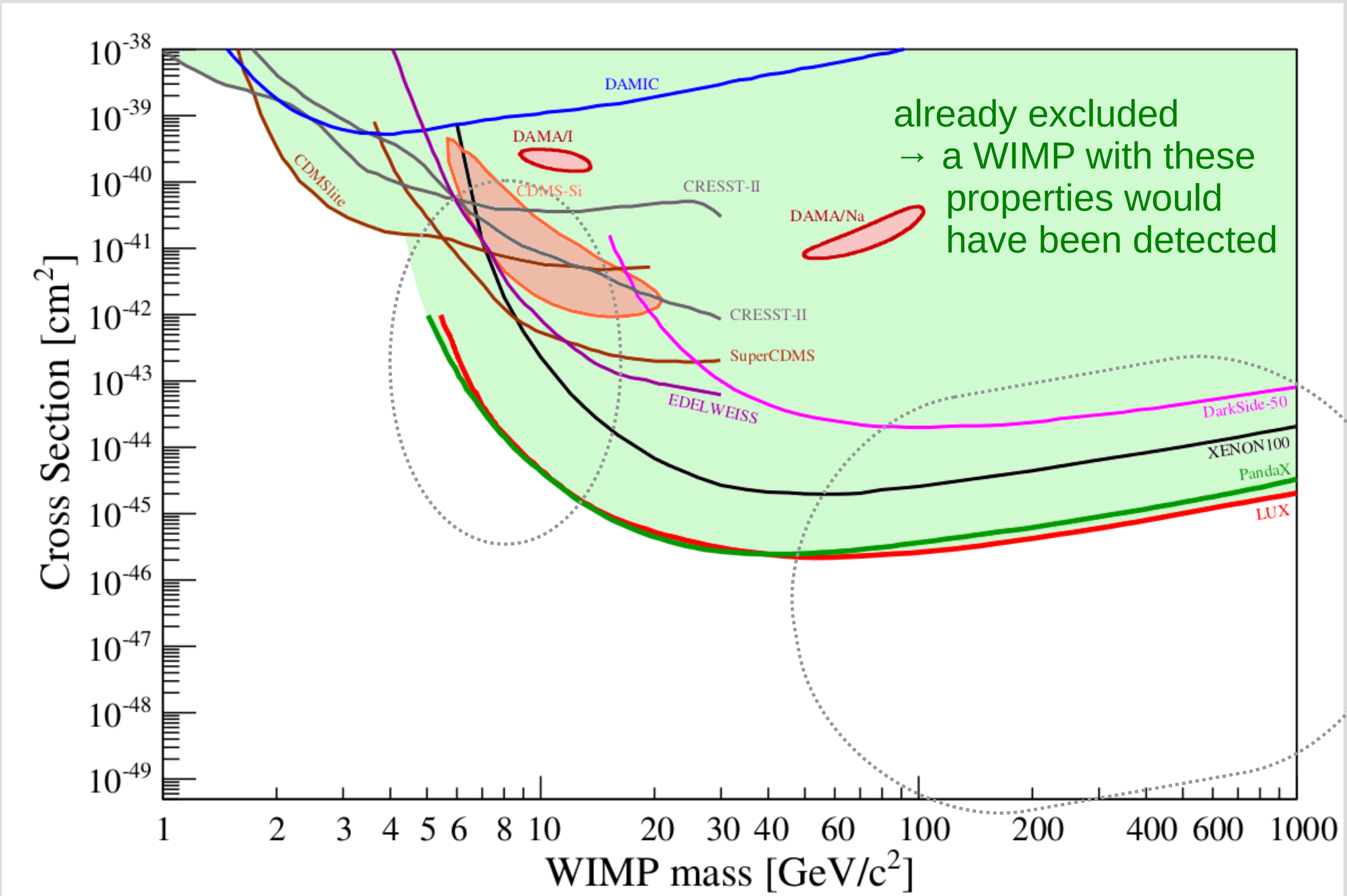
spin-independent WIMP-nucleon interactions



High WIMP-masses TPC dominated

→ ≥4.5 GeV/c²

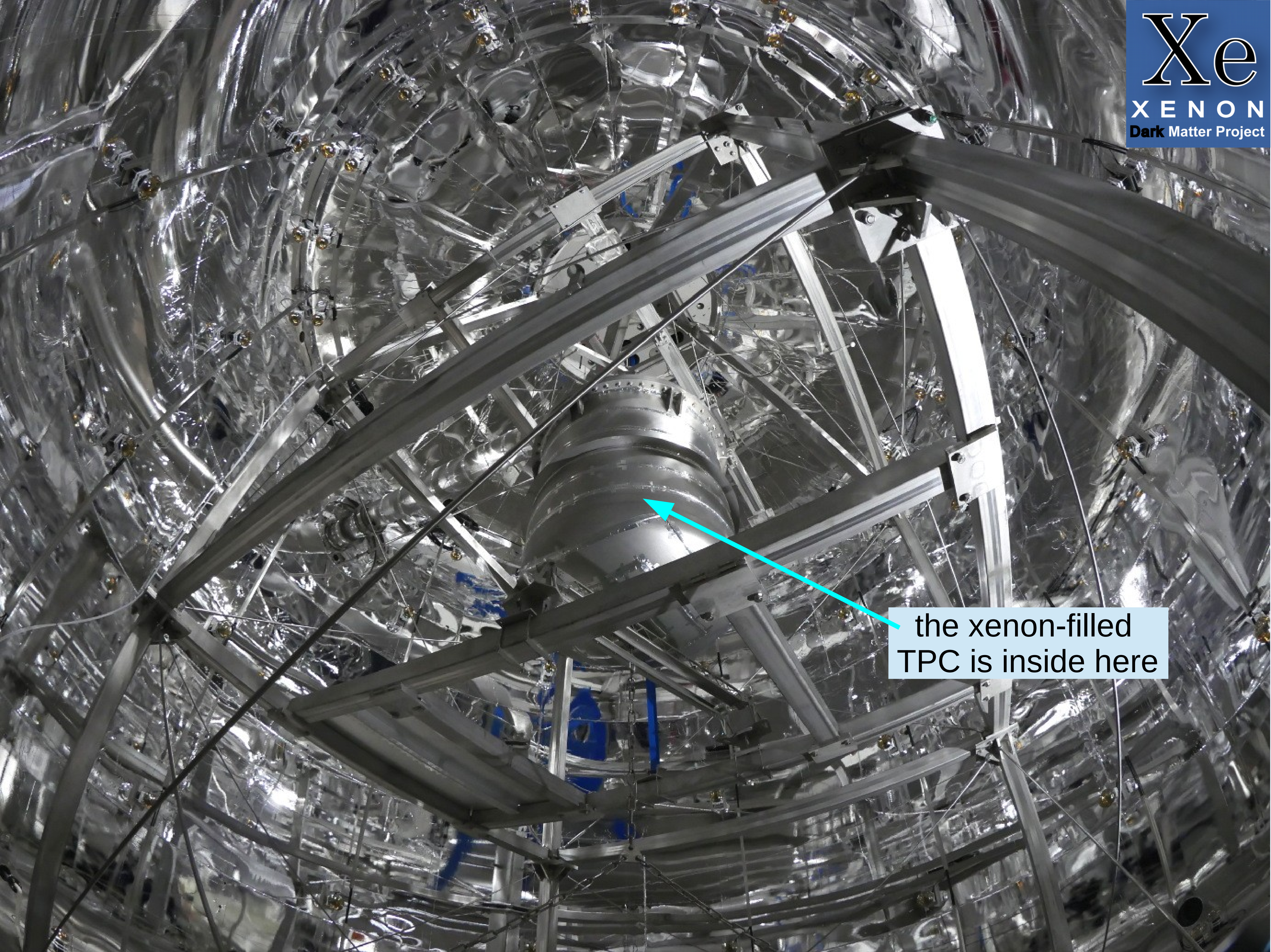
spin-independent WIMP-nucleon interactions



some projects are missing...

XENON1T @ LNGS

EPJ C 77, 881 (2017)


Xe
XENON
Dark Matter Project
the xenon-filled
TPC is inside here

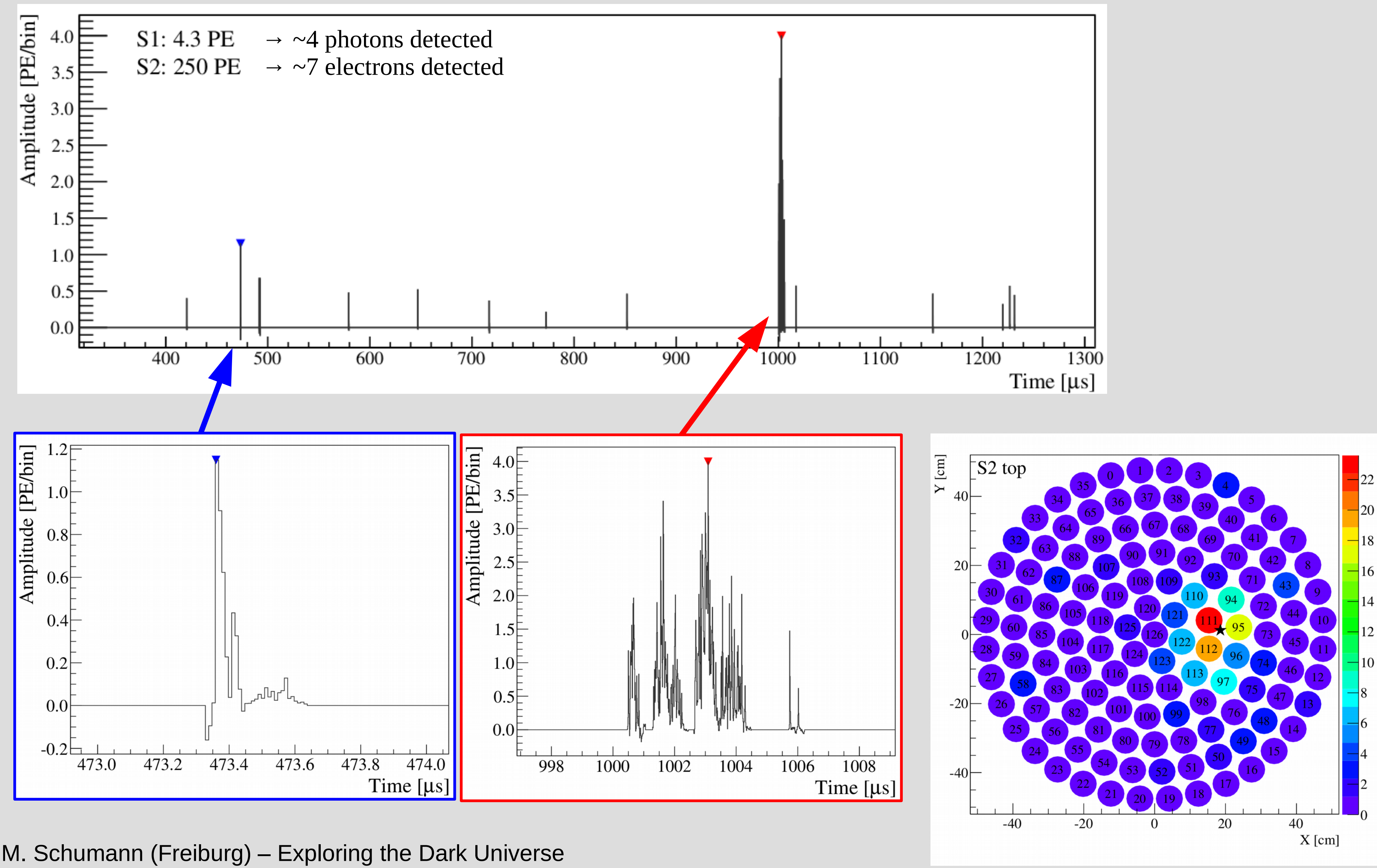
Xe
XENON
Dark Matter Project

largest LXe TPC ever built

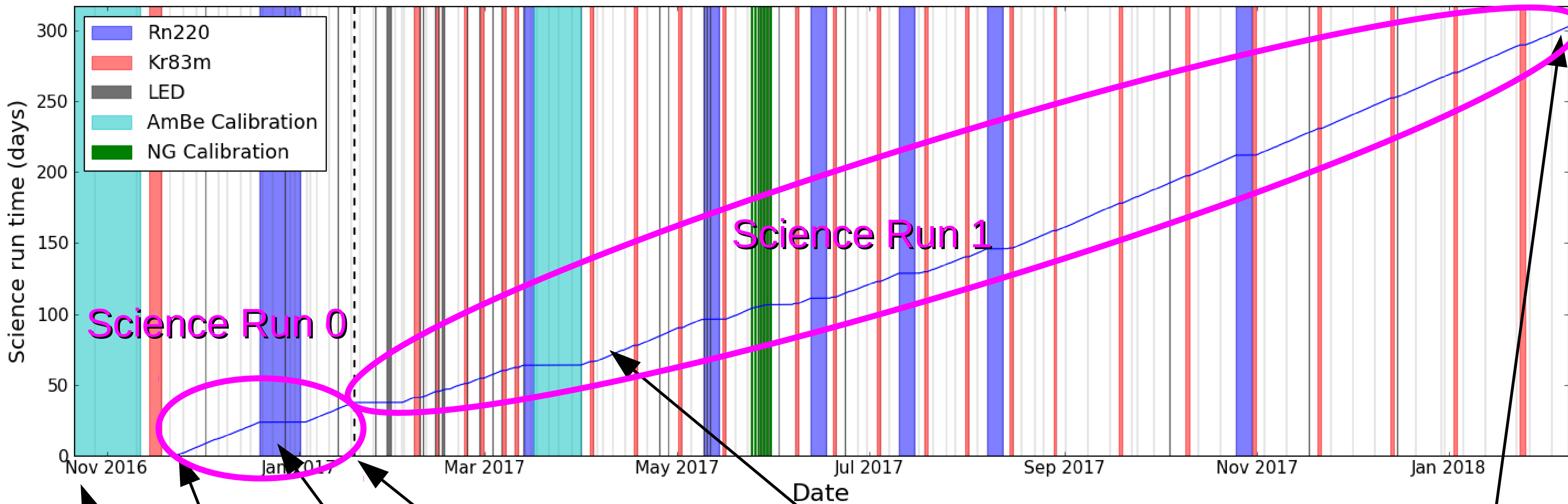
- **cylinder: 96 cm**
- **active LXe target: 2.0t** (3.2t total)
- **248 PMTs**

How would dark matter look?

… but it's a low-E neutron interaction from calibration!



Dark Matter Data Taking



commissioning completed

dark matter search starts

Xmas calibration

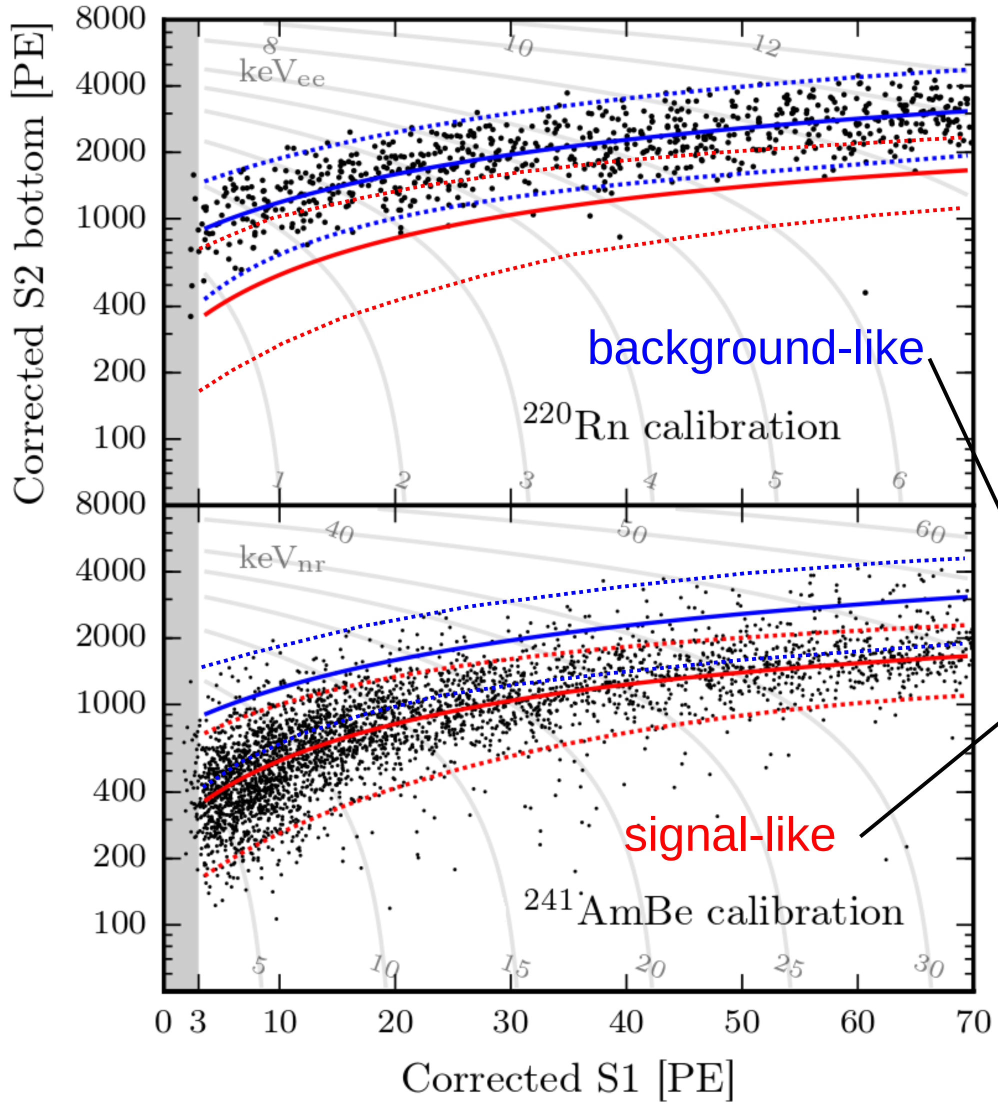
Earthquake!

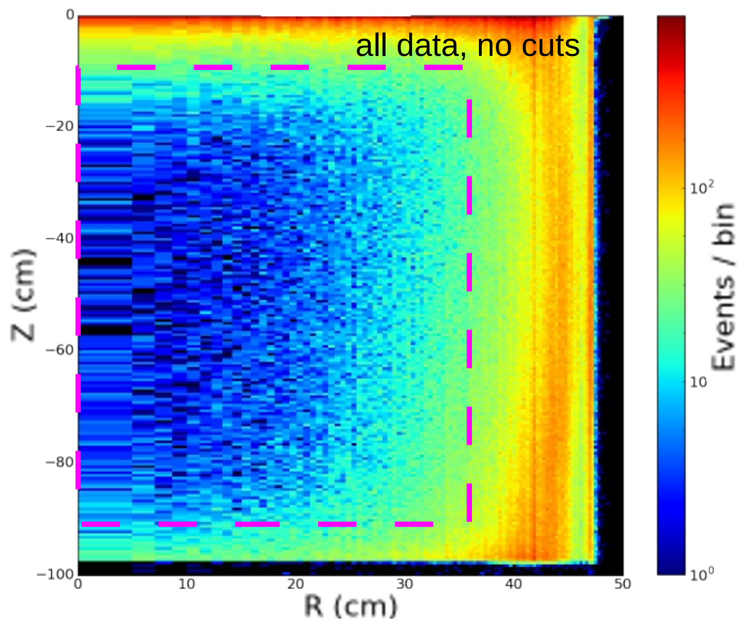
continue dark matter search and calibration

yesterday

Calibration and Analysis

PRL 119, 181301 (2017)





Used to construct background and signal models.

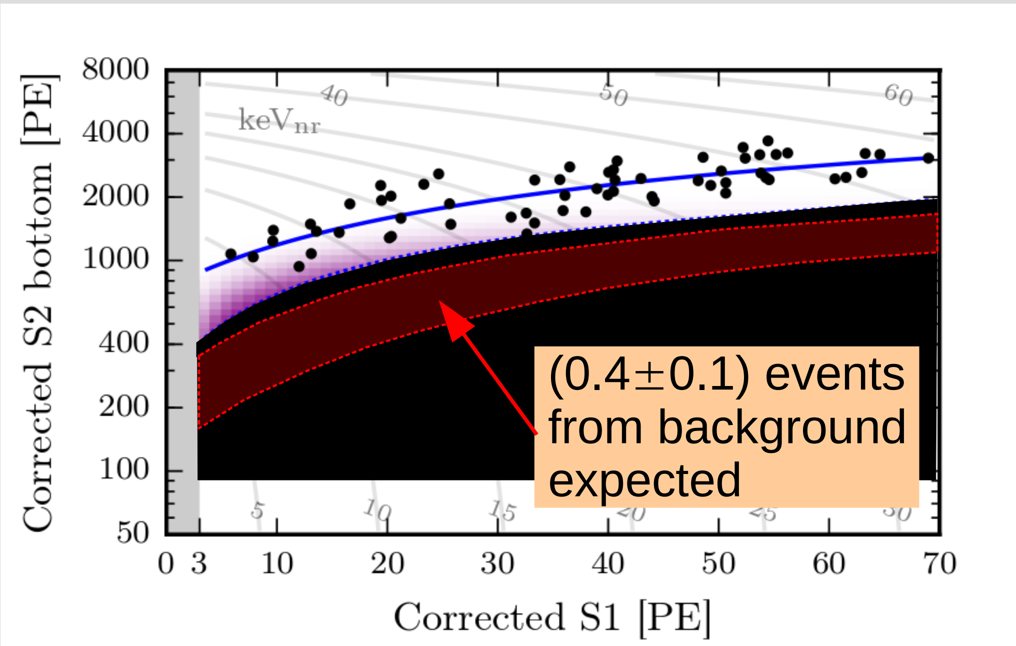
use **central 1000 kg** LXe for analysis

Blind analysis
= region of interest inaccessible during analysis to avoid human bias

Background Expectation

PRL 119, 181301 (2017)

figure of merit: exposure E = target mass [t] × measuring time [d]



XENON1T: **35.6 t × d**

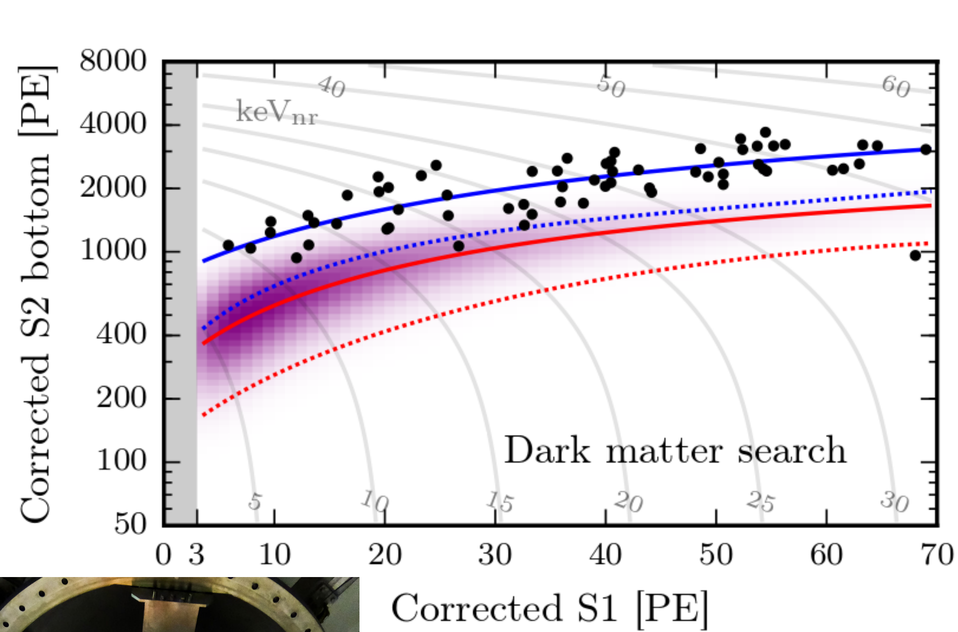
note:
final **unbinned profile likelihood** analysis takes into account

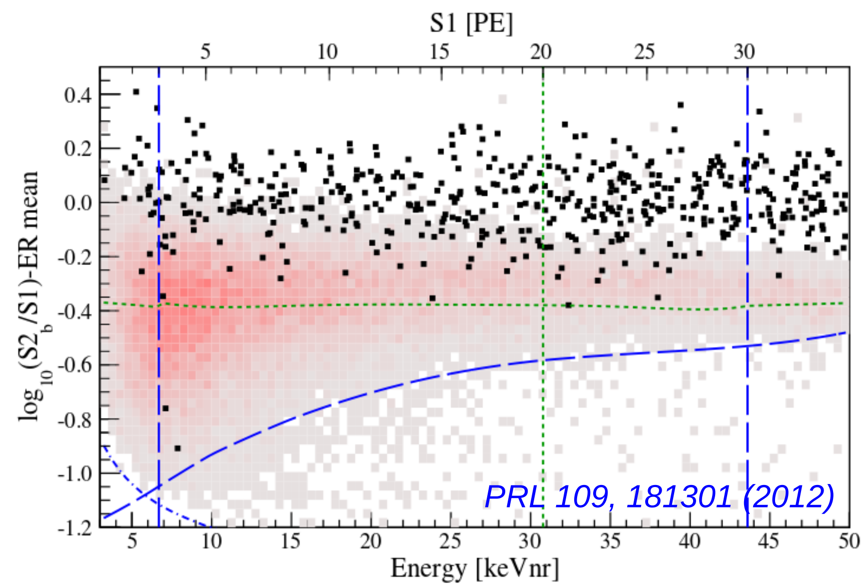
- full signal and background distributions
- full parameter space

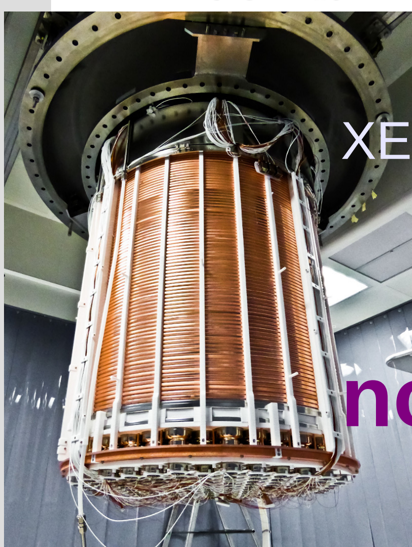
An ultra-low background

PRL 119, 181301 (2017)

figure of merit: exposure E = target mass [t] × measuring time [d]







XENON1T: **35.6 t × d** ← XENON100: **7.6 t × d**

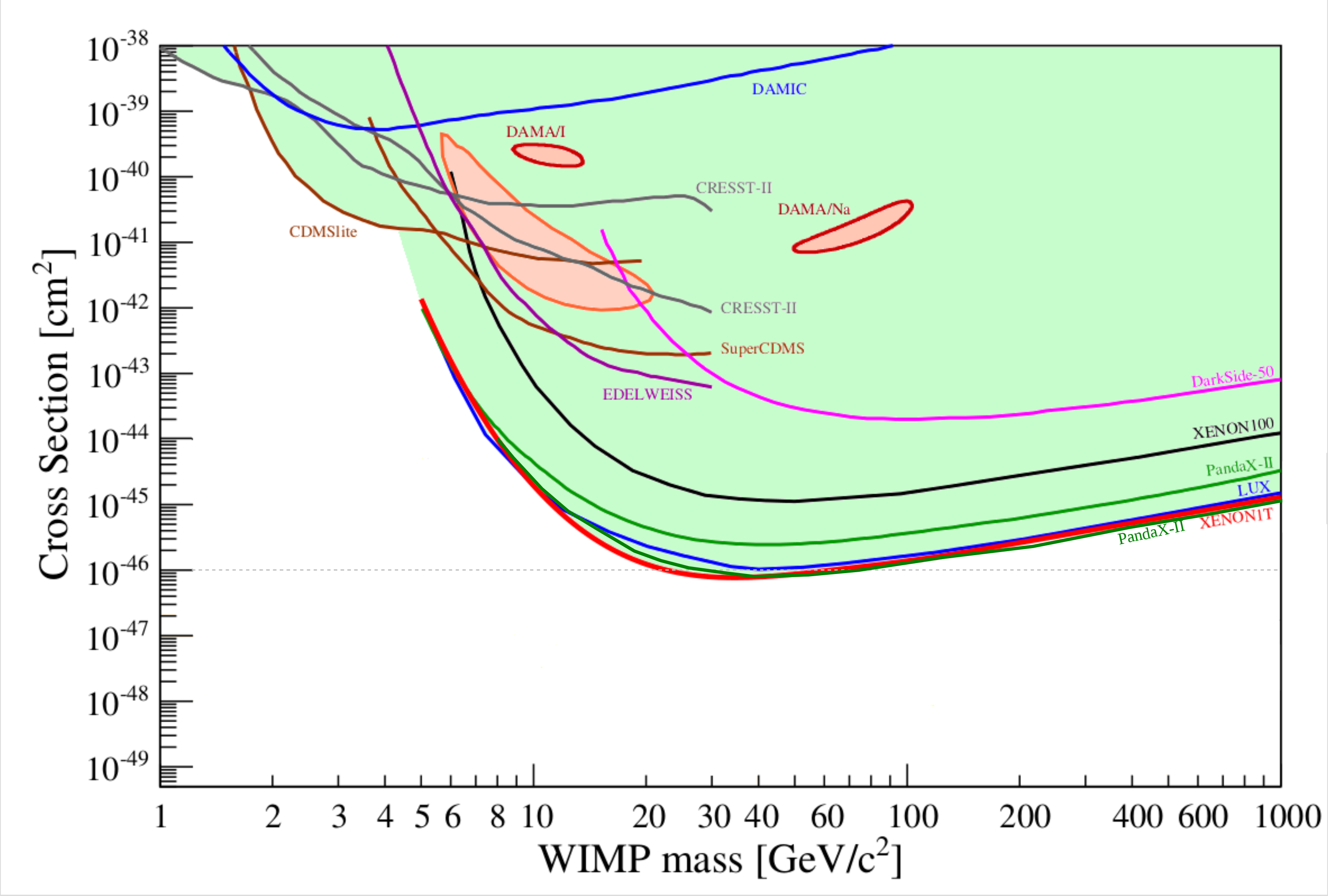
5× more data
30× lower background

no dark matter candidate observed!

No Signal → Exclusion Limit

PRL 119, 181301 (2017)

spin-independent WIMP-nucleon interactions



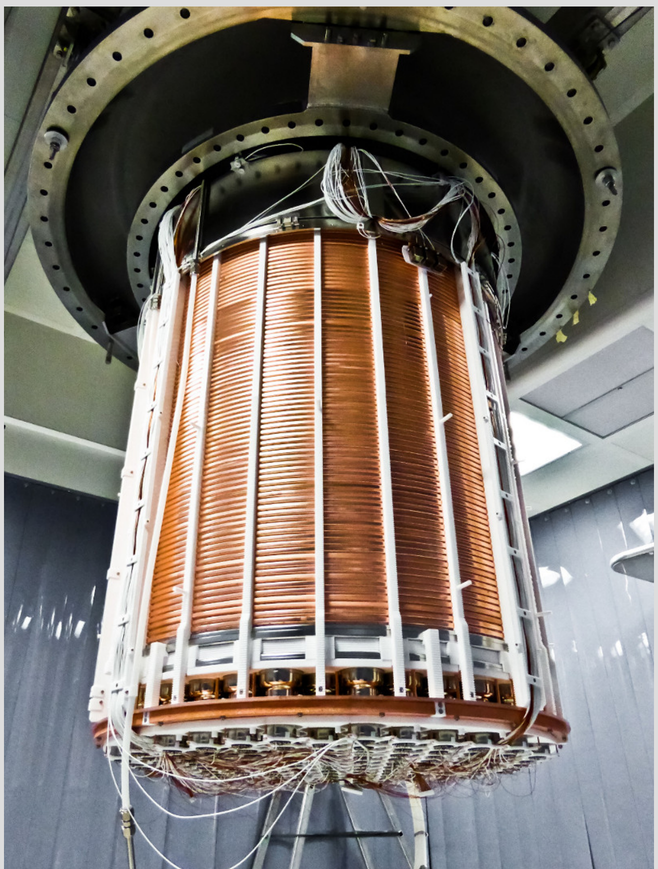
XENON1T science goal: 5× more sensitive than current result

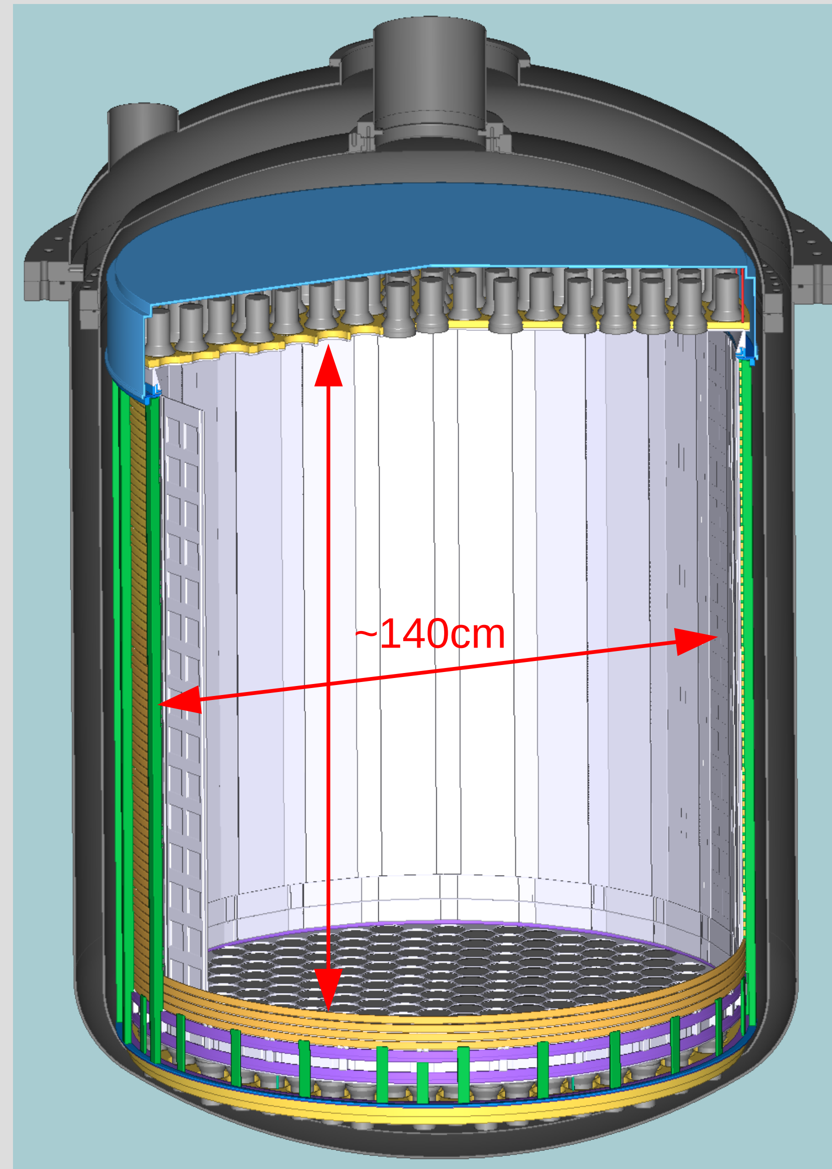
XENON1T → XENONnT

JCAP 04, 027 (2016)

XENON1T
- **2t** active LXe target
- taking data

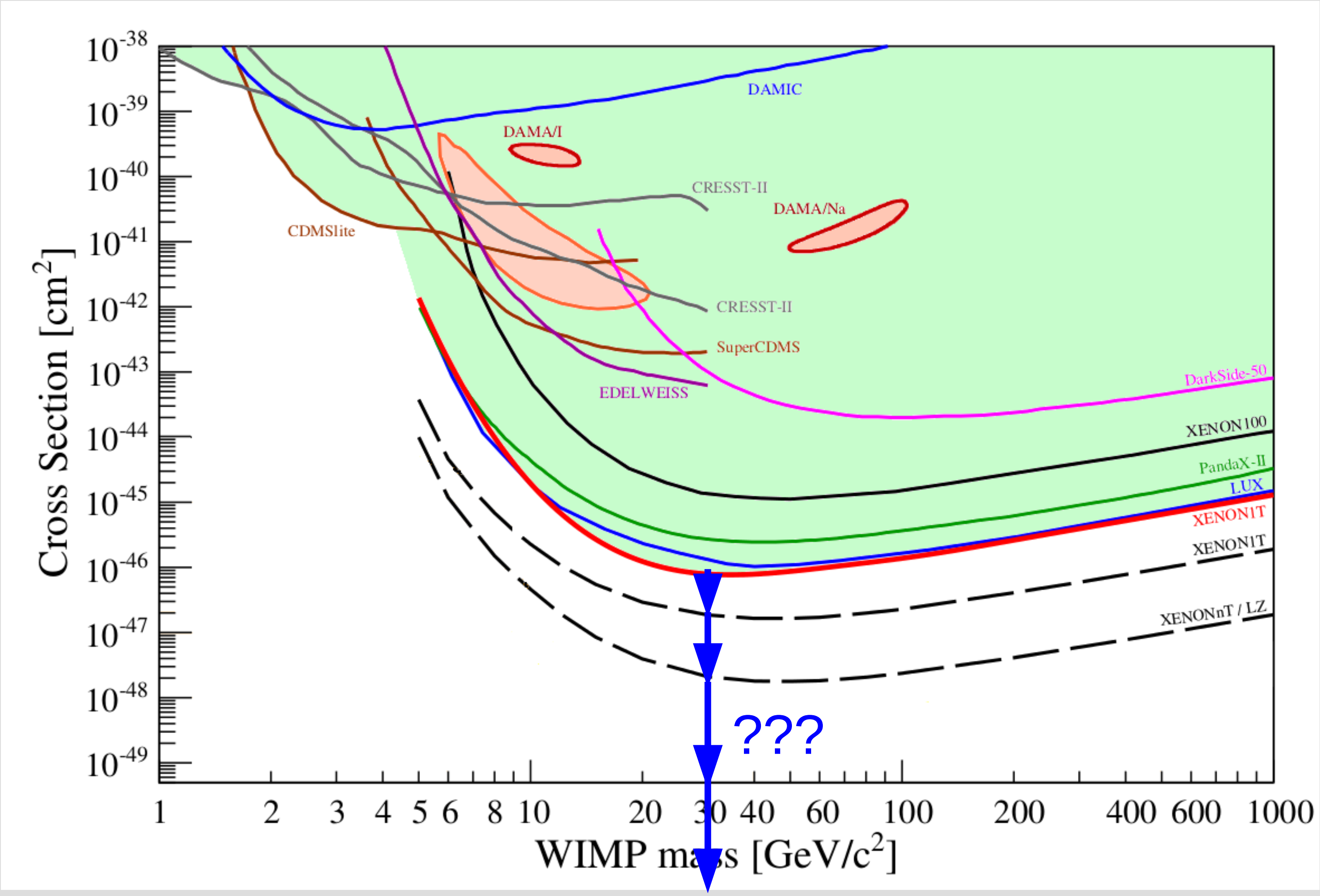
XENONnT
- **6t** active target
- science run by 2019





XENON: The Future

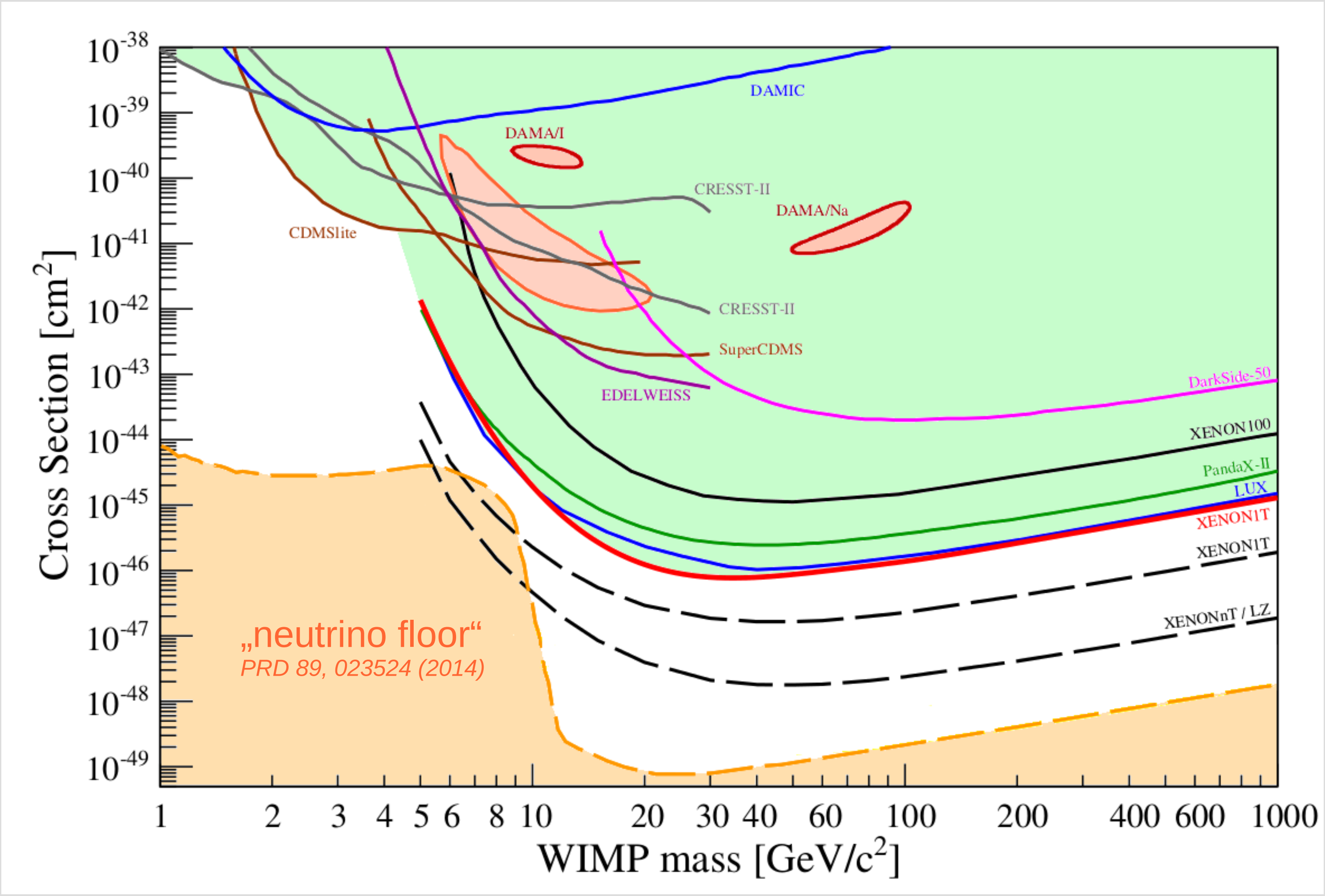
spin-independent WIMP-nucleon interactions



some projects are missing...

Dark Matter Searches: The Limit

spin-independent WIMP-nucleon interactions



some projects are missing...

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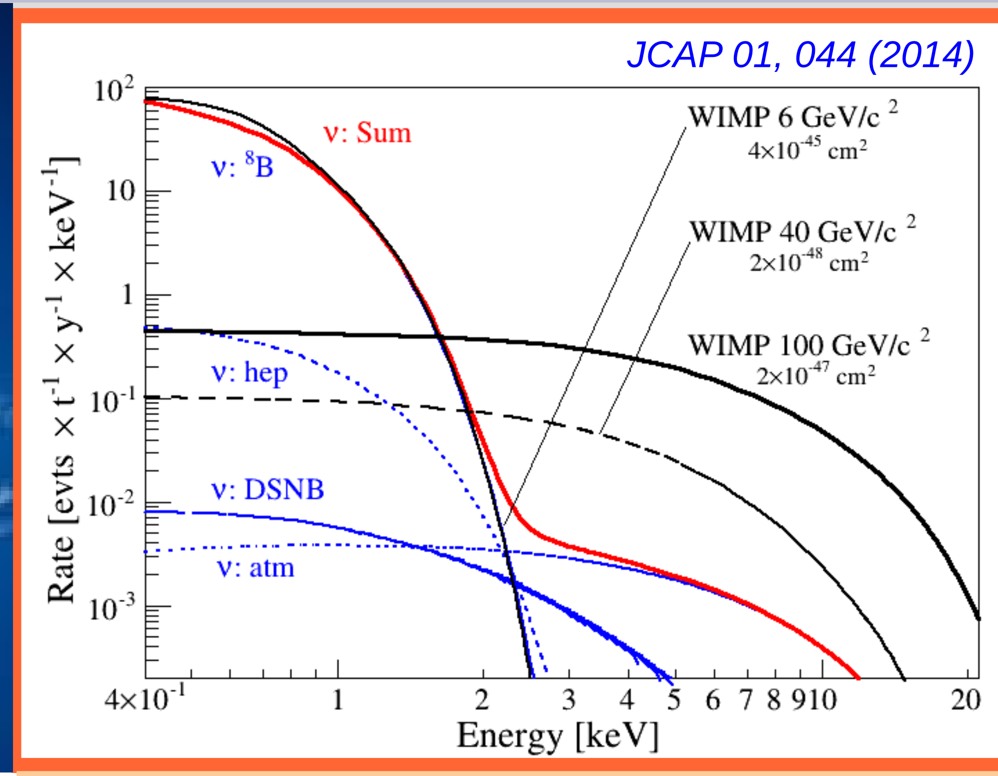
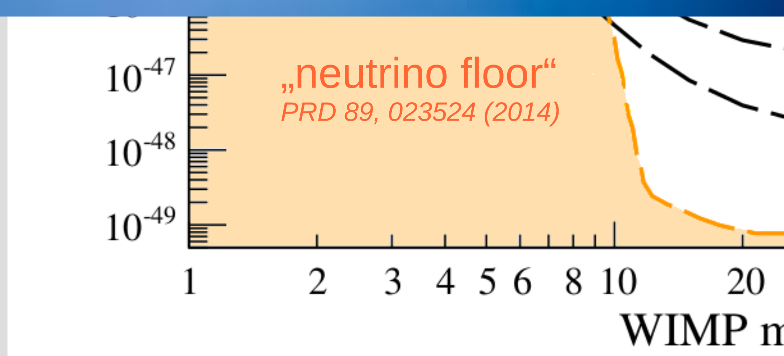
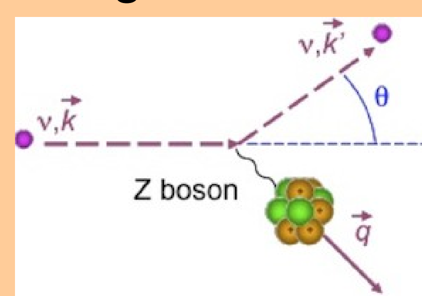


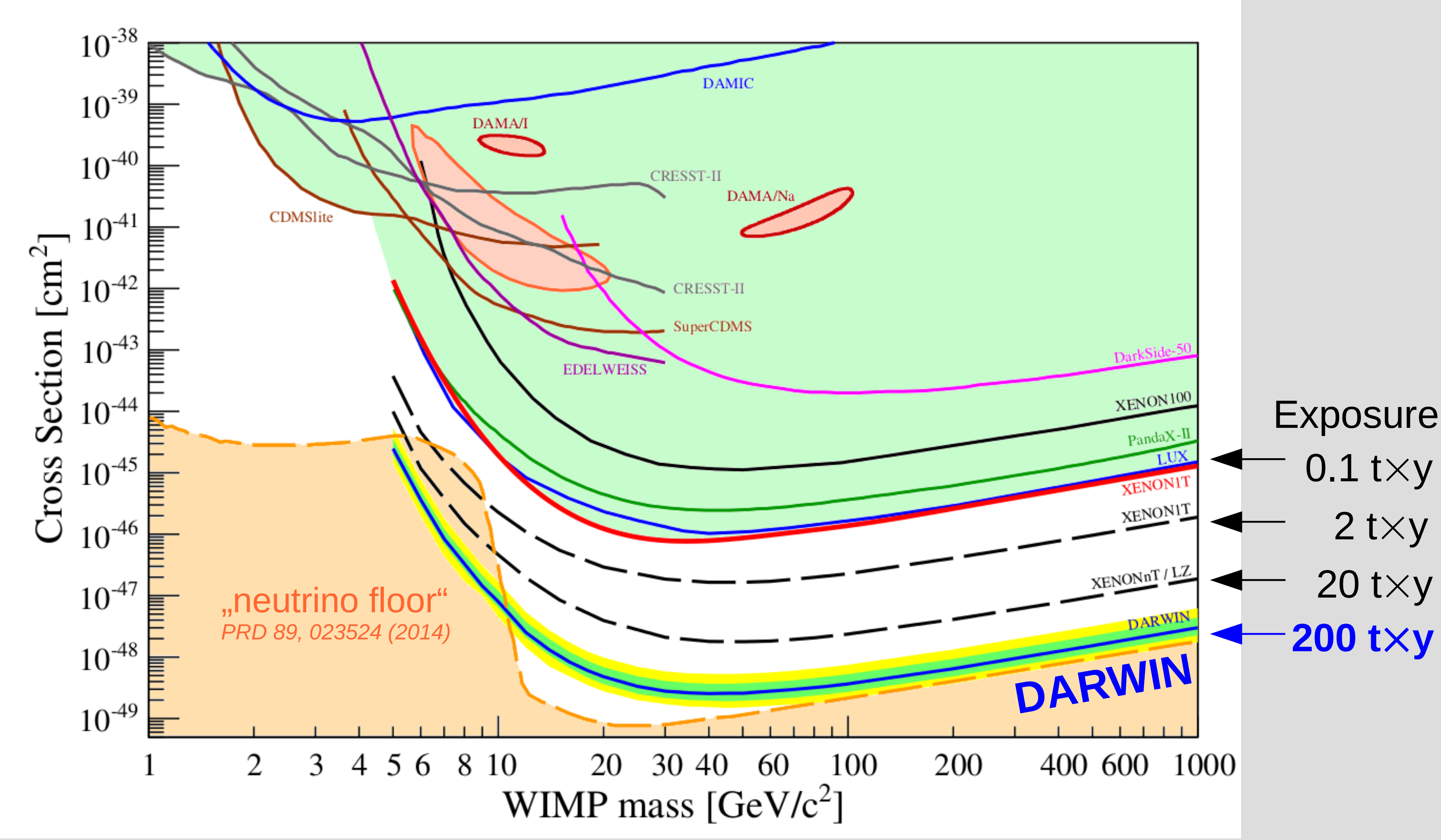


Interactions from coherent neutrino-nucleus scattering (CNNS) will dominate
→ **ultimate background** for direct detection



DARWIN The ultimate WIMP Detector

spin-independent WIMP-nucleon interactions

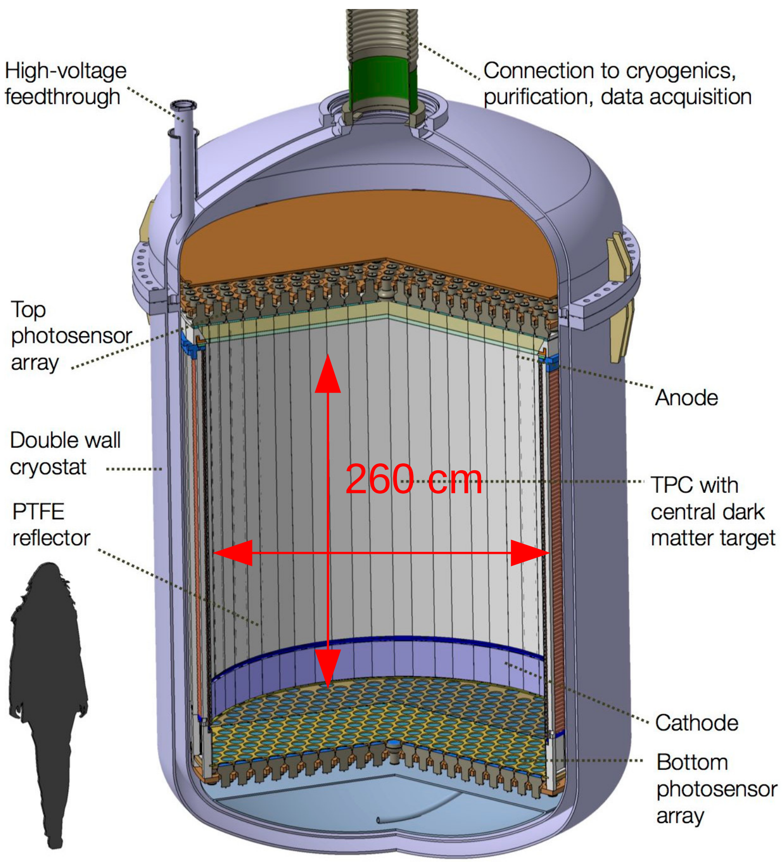


some projects are missing...

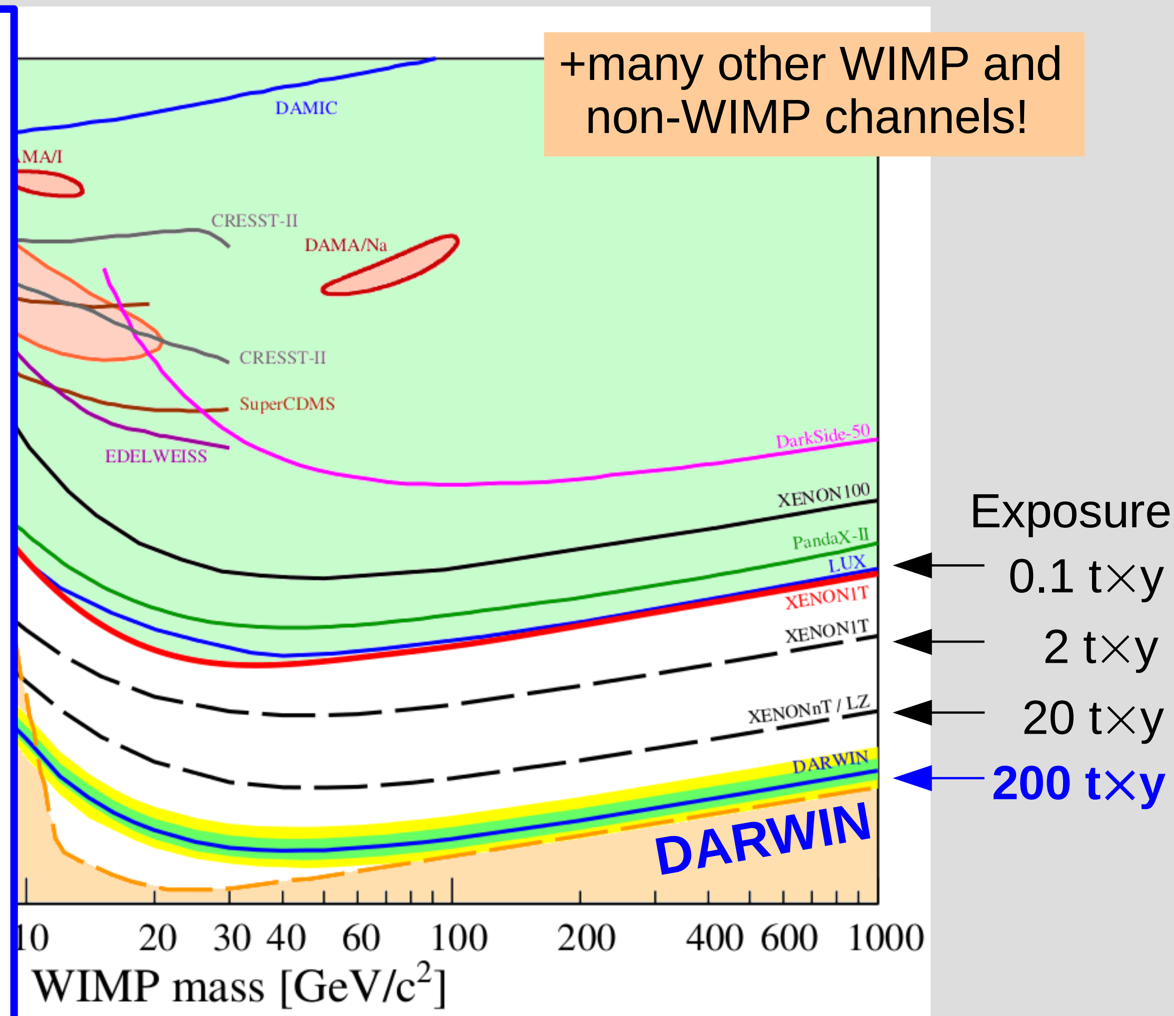
DARWIN The **ultimate** WIMP Detector

JCAP 11, 017 (2016)

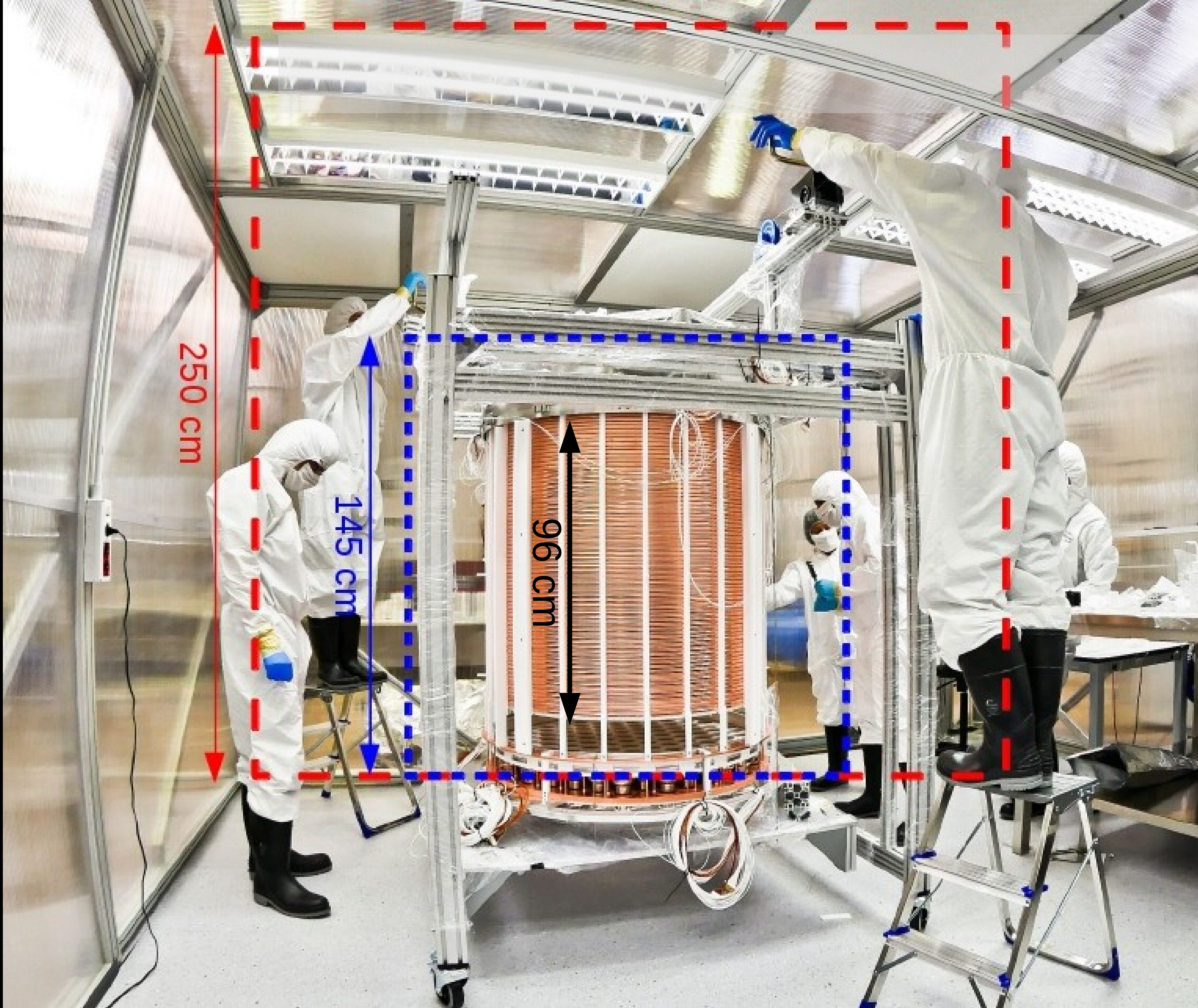
Baseline scenario
~50t total LXe mass
~40 t LXe TPC
~30 t fiducial mass



spin-independent WIMP-nucleon interactions

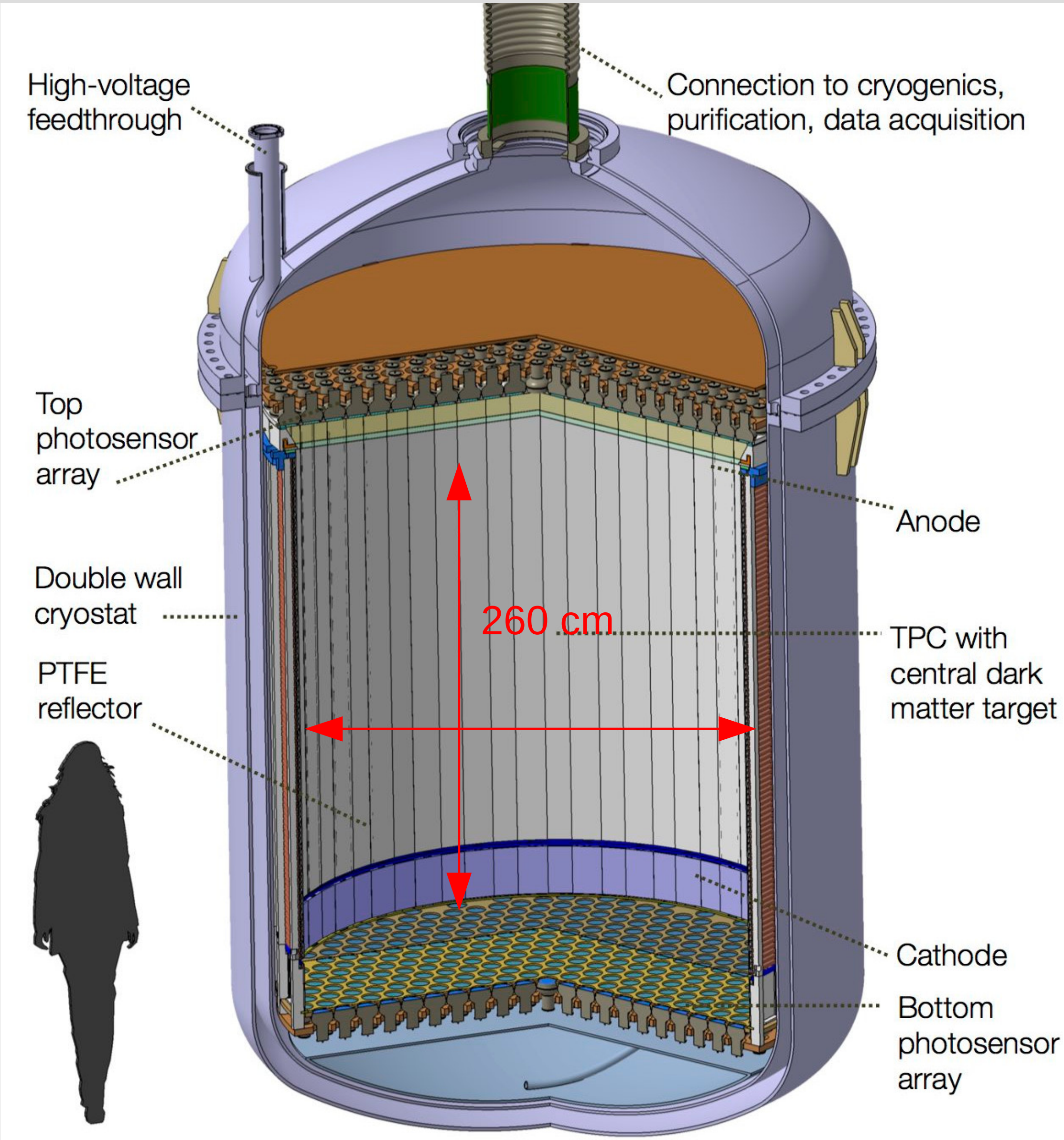


some projects are missing...


250 cm
145 cm
96 cm

DARWIN The ultimate WIMP Detector

JCAP 11, 017 (2016)

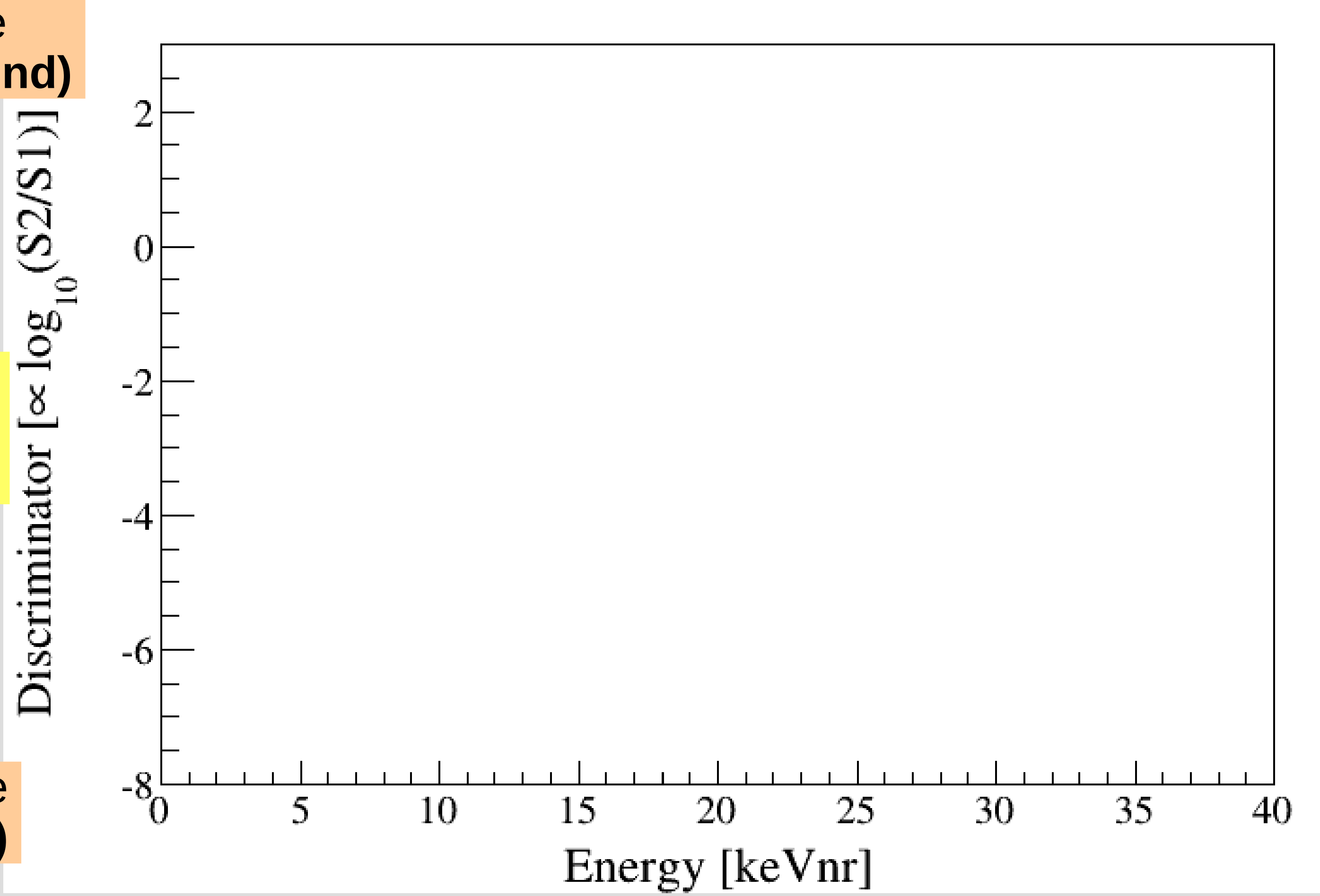


Challenges

- Size
  - → electron drift (HV)
  - → diameter (TPC electrodes)
  - → mass (LXe purification)
  - → dimensions (radioactivity)
  - → detector response (calibration, corrections)
- Backgrounds
  - → ^{222}Rn: factor 100 required
  - → (α,n) neutrons (from PTFE)
- Photosensors
  - → high light yield (QE)
  - → low radioactivity
  - → long-term stability
- etc etc

ULTIMATE

WIMP Detection



ER-like
(background)

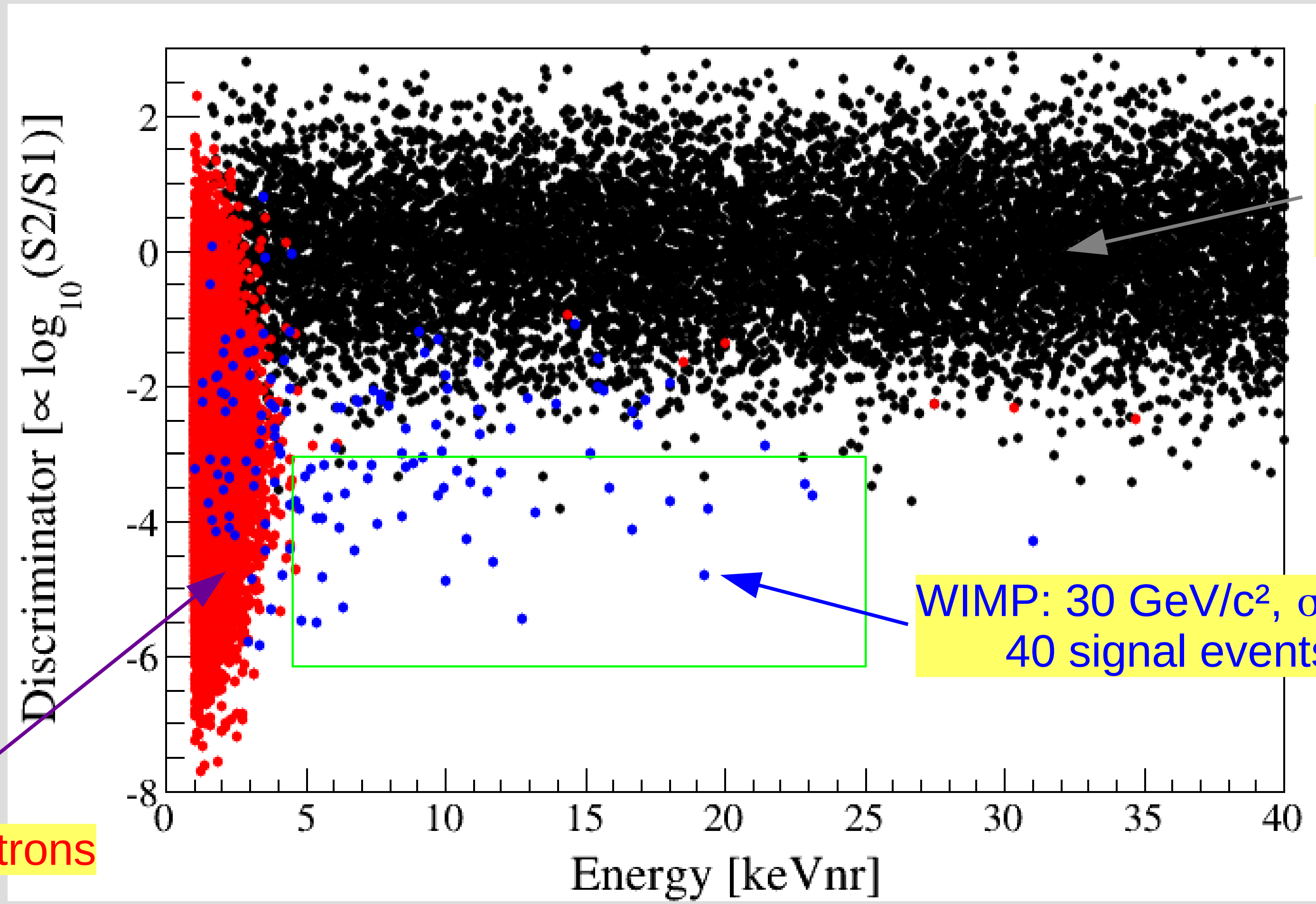
discrimination variable;
based on S2/S1

NR-like
(signal)

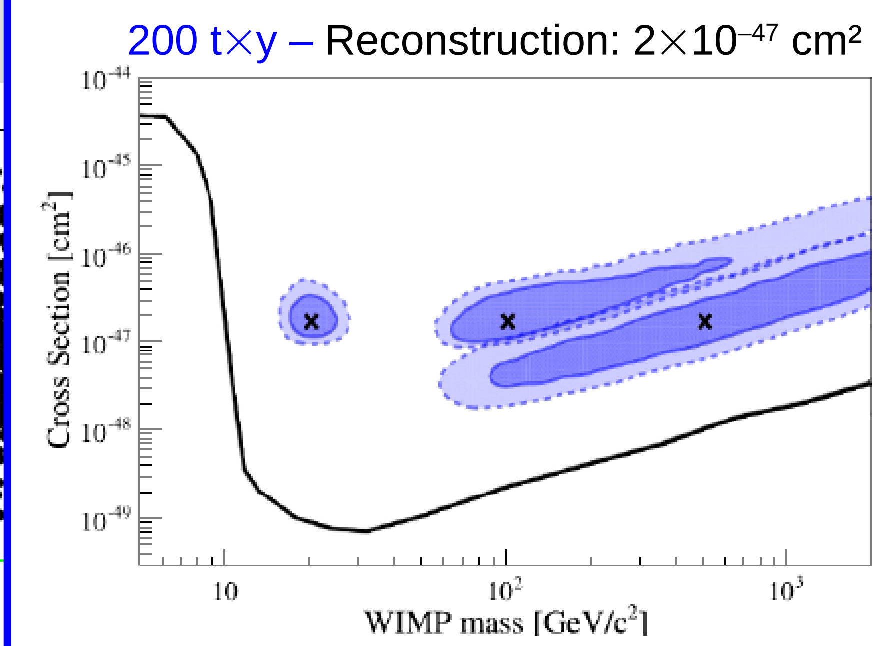
reconstructed energy,
based on S1 and S2 signal

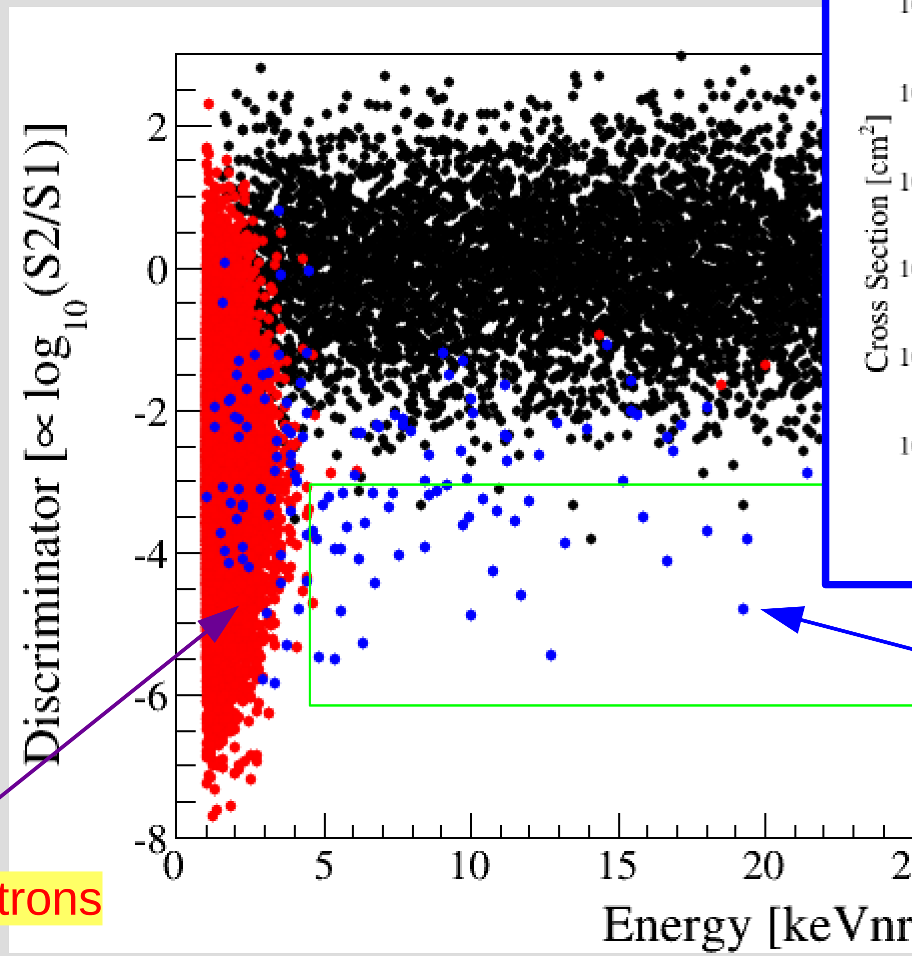
WIMP Detection

DARWIN



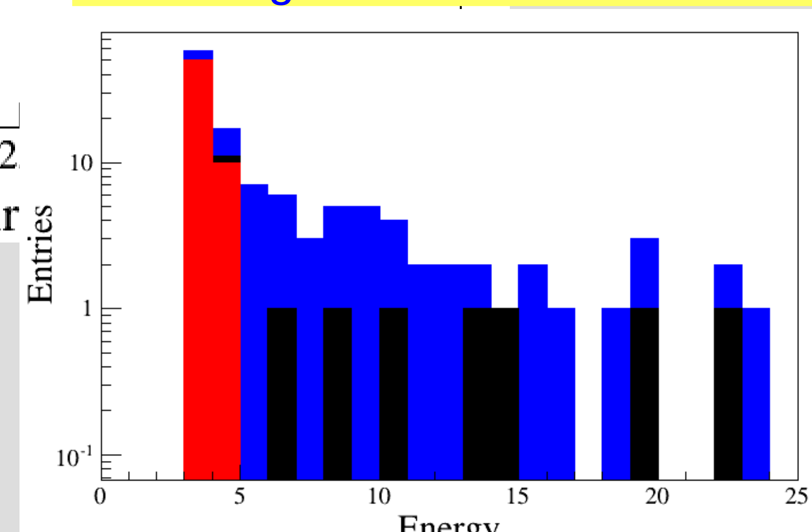
WIMP Spectroscopy



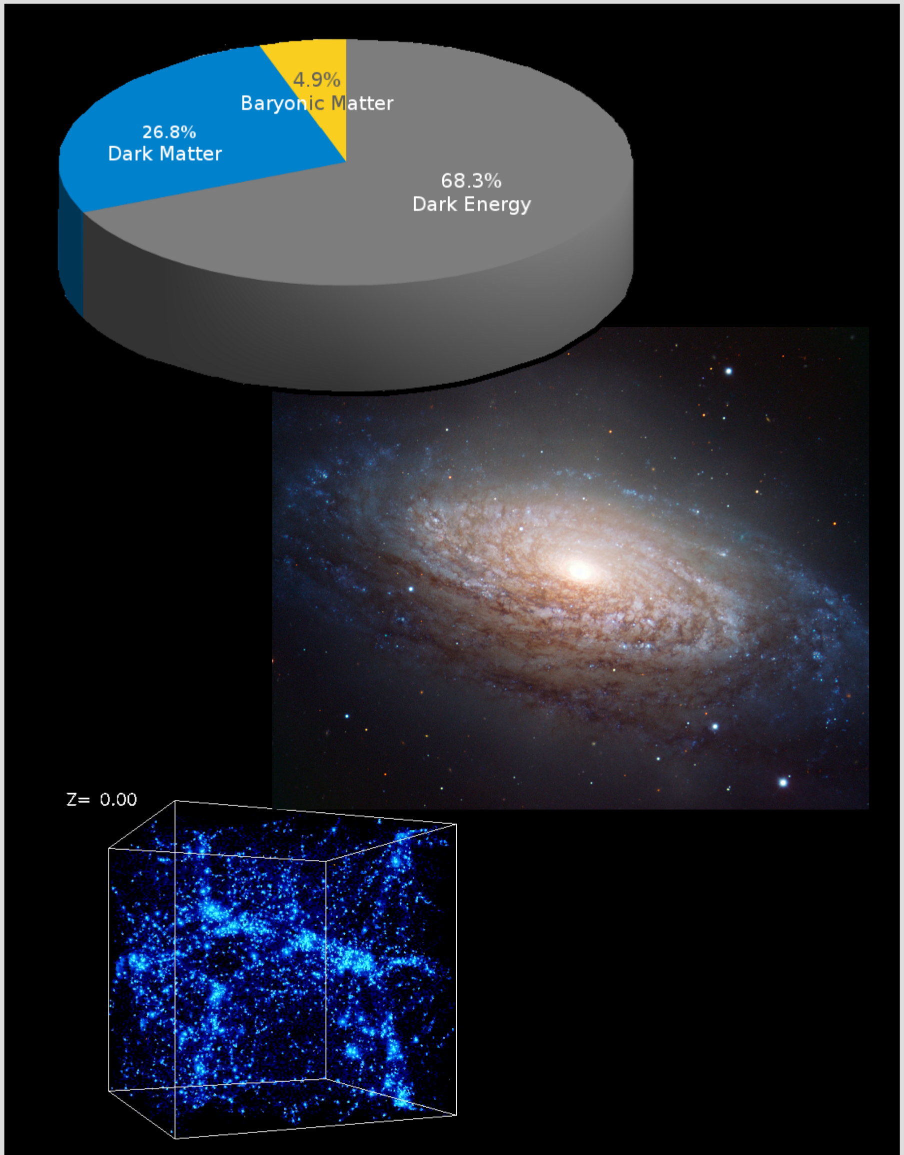


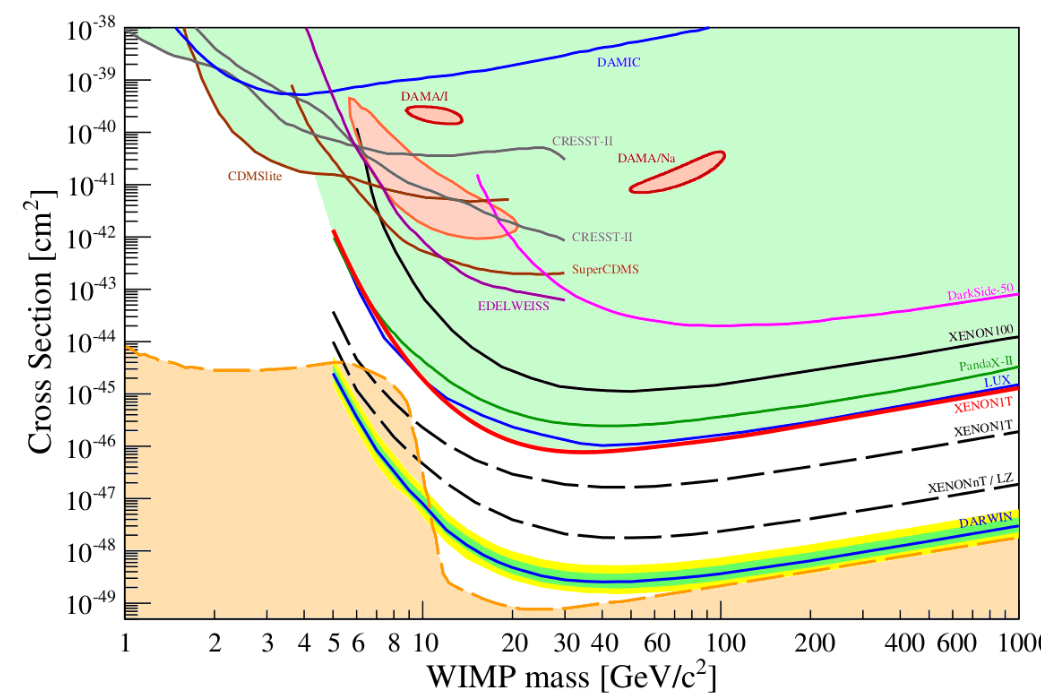
CNNS+neutrons

WIMP: 30 GeV/c², σ=2×10⁻⁴⁸ cm²
40 signal events in box



Exploring the dark with LXe Detectors





www.xenon1t.org

www.darwin-observatory.org